Limits on Lorentz Invariance Violation from Fermi Gamma-Ray Space Telescope Observations of Gamma-Ray Bursts

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on behalf of the Fermi LAT & GBM Collaborations

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Outline of the Talk:

- Brief motivation & parameterization
- Constraining Lorentz Invariance Violation using GRBs: why use GRBs & how do we set limits
- Limit from the bright long GRB 080916C at $z \sim 4.35$
- 3 different types of limits from the short bright GRB 090510 at $z = 0.903$: detailed description & results
- **Summary** of limits on LIV using Fermi LAT GRBs
- **Prospects for CTA** LIV studies using GRBs
- **Conclusions**

Brief Motivation & Parameterization

- Some quantum-gravity (QG) models allow or even predict (e.g. Ellis et al. 2008) Lorentz invariance violation (LIV)
- We directly constrain a simple form of LIV - dependence of the speed of light on the photon energy: $v_{ph}(E_{ph}) \neq c$
- This may be parameterized through a Taylor expansion of the LIV terms in the dispersion relation:

$$c^2 p_{ph}^2 = E_{\mathrm{ph}}^2 \left[1 + \sum_{k=1}^{\infty} s_k \left(\frac{E_{\mathrm{ph}}}{\mathrm{M}_{\mathrm{QG,k}} c^2} \right)^k \right] \quad , \quad \text{where } \mathrm{M}_{\mathrm{QG,k}} \leq \mathrm{M}_{\mathrm{Planck}} \text{ is naturally expected}$$

- $s_k \in \{-1, 0, 1\}$ stresses the model dependent sign of the effect
- The most natural scale for LIV is the **Planck scale**, where quantum effects on the structure of space-time are expected to become strong: $l_{\mathrm{Planck}} = (\hbar G/c^3)^{1/2} \approx 1.62\times10^{-33}$ cm

$$E_{\mathrm{Planck}} = M_{\mathrm{Planck}} c^2 = (\hbar c^5/G)^{1/2} \approx 1.22\times10^{19}\ \mathrm{GeV}$$

Brief Motivation & Parameterization

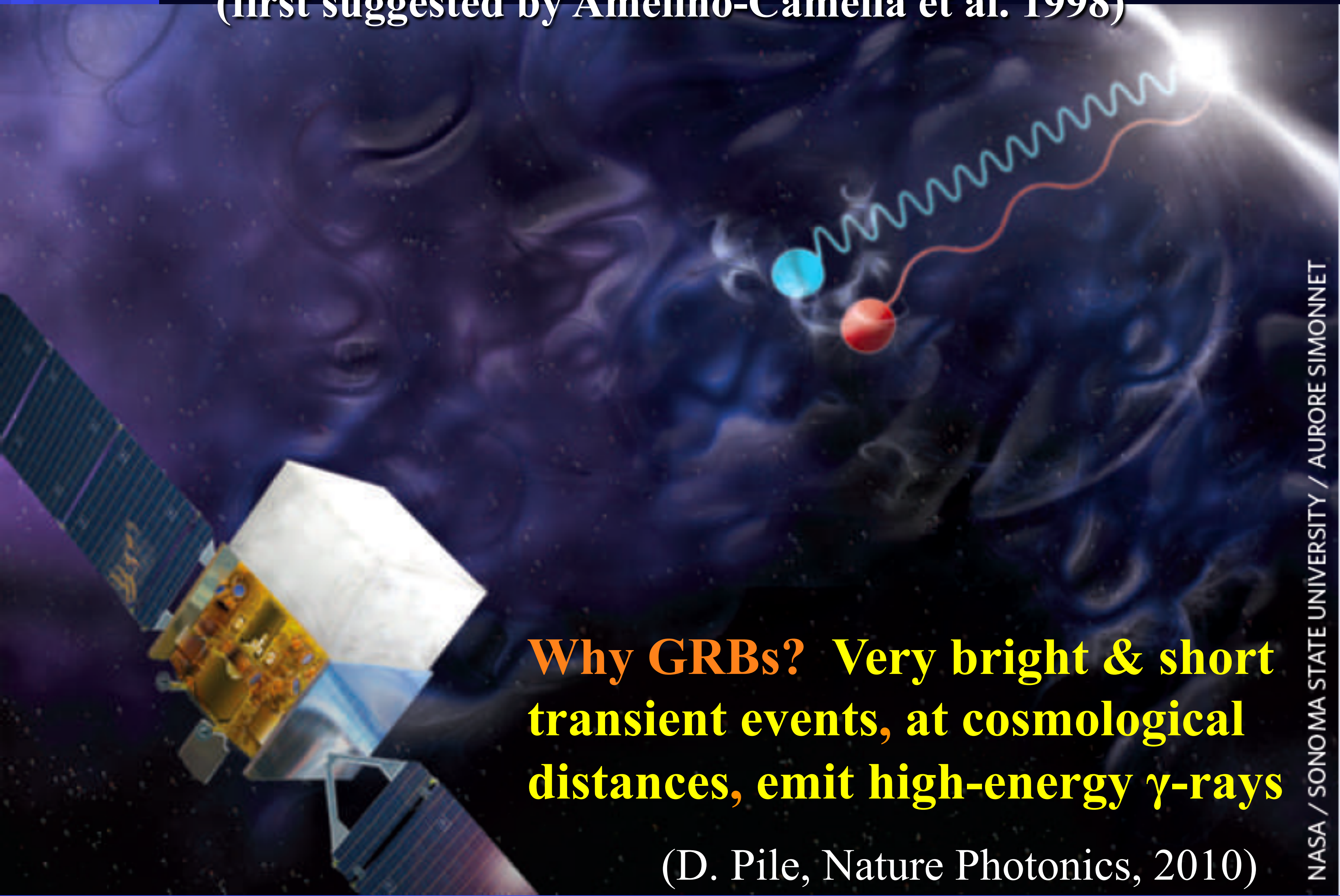
- The photon propagation speed is given by the group velocity:

$$c^2 p_{ph}^2 = E_{\mathrm{ph}}^2 \left[1 + \sum_{k=1}^{\infty} s_k \left(\frac{E_{\mathrm{ph}}}{\mathrm{M}_{\mathrm{QG,k}} c^2} \right)^k \right] \quad , \quad v_{\mathrm{ph}} = \frac{\partial E_{\mathrm{ph}}}{\partial p_{\mathrm{ph}}} \approx c \left[1 - s_n \frac{(1+n)}{2} \left(\frac{E_{\mathrm{ph}}}{\mathrm{M}_{\mathrm{QG,n}} c^2} \right)^n \right]$$

- Since $E_{ph} \ll M_{QG,k}c^2 \lesssim E_{planck} \sim 10^{19}$ GeV the **lowest order** non-zero term, of order $n = \min\{k \mid s_k \neq 0\}$, **dominates**
- Usually $n = 1$ (linear) or 2 (quadratic) are considered
- We focus here on $\mathbf{n = 1}$, since only in this case are our limits of the order of the Planck scale
- We try to constrain **both possible signs** of the effect:
  - $s_n = 1$, $v_{ph} < c$: higher energy photons propagate slower
  - $s_n = -1$, $v_{ph} > c$: higher energy photons propagate faster
- We stress: here $c = v_{ph}(E_{ph} \to 0)$ is the low energy limit of v_{ph}

Constraining LIV Using GRBs

(first suggested by Amelino-Camelia et al. 1998)



Why GRBs? Very bright & short transient events, at cosmological distances, emit high-energy γ-rays

(D. Pile, Nature Photonics, 2010)

Constraining LIV Using GRBs

- A high-energy photon E_h would arrive after (in the sub-luminal case: $v_{ph} < c$, $s_n = 1$), or possibly before (in the super-luminal case, $v_{ph} > c$, $s_n = -1$) a low-energy photon E_l emitted together
- The time delay in the arrival of the high-energy photon is:

$$\Delta t_{\mathrm{LIV}} = s_n \frac{(1+n)}{2H_0} \frac{E_h^n - E_l^n}{\left(\mathrm{M}_{\mathrm{QG,n}} c^2\right)^n} \int_0^z \frac{(1+z')^n}{\sqrt{\Omega_m (1+z')^3 + \Omega_\Lambda}} dz'$$

(Jacob & Piran 2008)

- The photons E_h & E_l do not have to be emitted at exactly the same time & place in the source, but we must be able to limit the difference in their effective emission times (i.e. in the times they would arrive to us without any energy dispersion)
- Our limits apply to any source of energy dispersion on the way from the source to us, and may constrain some (even more) exotic physics ($\Delta t_{\mathrm{LIV}} \rightarrow \Delta t_{\mathrm{LIV}} + \Delta t_{\mathrm{exotic}}$)

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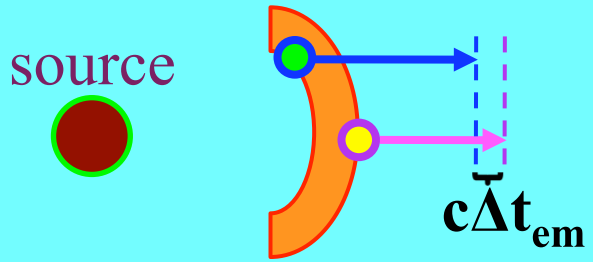


$$\Delta t_{obs} = \Delta t_{em} + \Delta t_{LIV}$$

Method 1

- Limits only $s_n = 1$ - the sub-luminal case: $v_{ph} < c$, & positive time delay, $\Delta t_{LIV} = t_h - t_{em} > 0$ (here t_h is the actual measured arrival time, while t_{em} would be the arrival time if $v_{ph} = c$)
- We consider a single high-energy photon of energy E_h and assume that it was emitted after the onset time (t_{start}) of the relevant low-energy (E_l) emission episode: $t_{em} > t_{start}$
- $\Rightarrow \Delta t_{LIV} = t_h - t_{em} < t_h - t_{start}$
- A conservative assumption: t_{start} = the onset of any observed emission from the GRB

Limits on LIV: GRB080916C ($z \approx 4.35$)

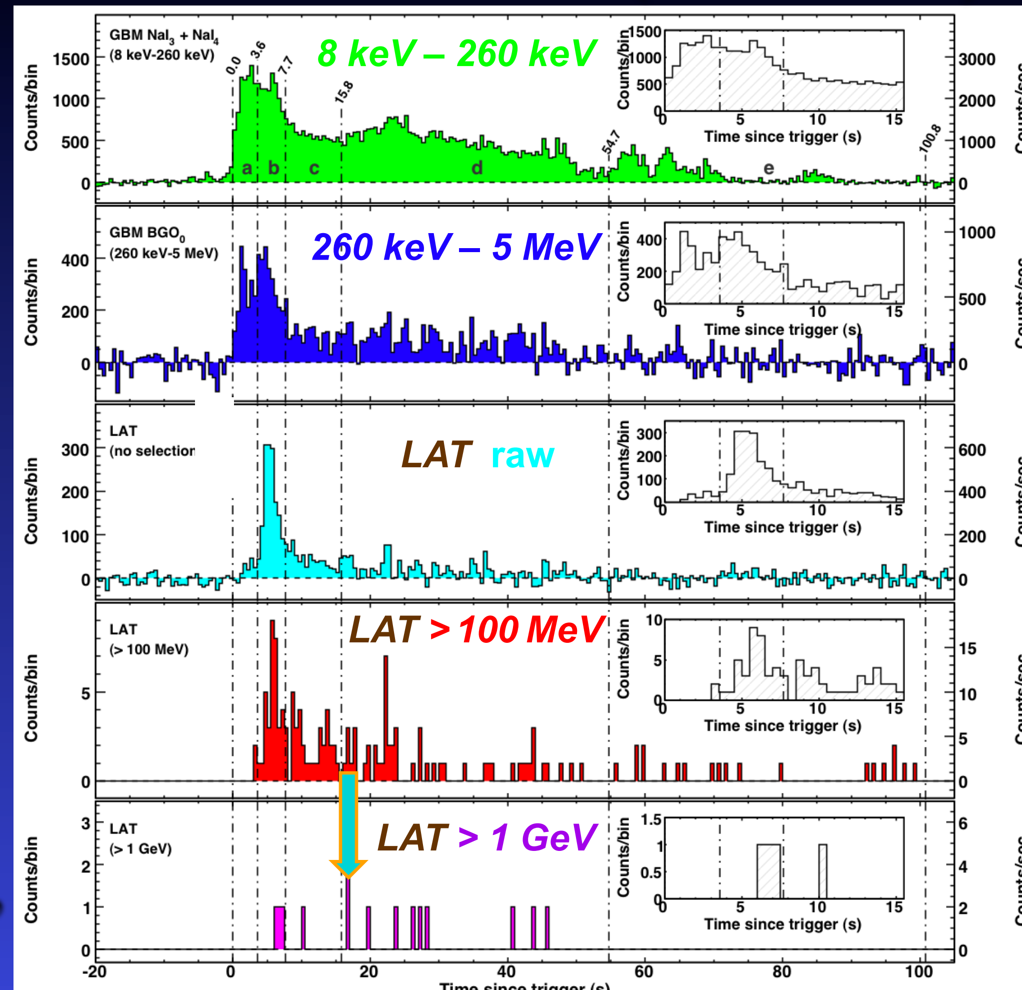
- GRB080916C: highest energy photon (13 GeV) arrived 16.5 s after low-energy photons started arriving (=the GRB trigger)

  $\Rightarrow$ **conservative lower limit:**

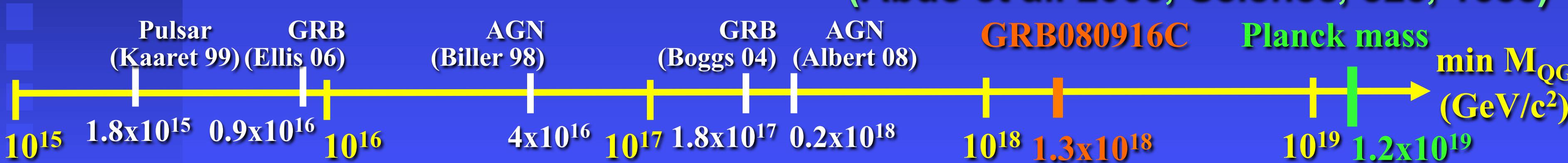
  $\mathbf{M_{QG,1} > 1.3\times10^{18}\ GeV/c^2}$

  $\approx 0.11 M_{Planck}$

- This improved upon the previous limits of this type, reaching 10% of M_{Planck}

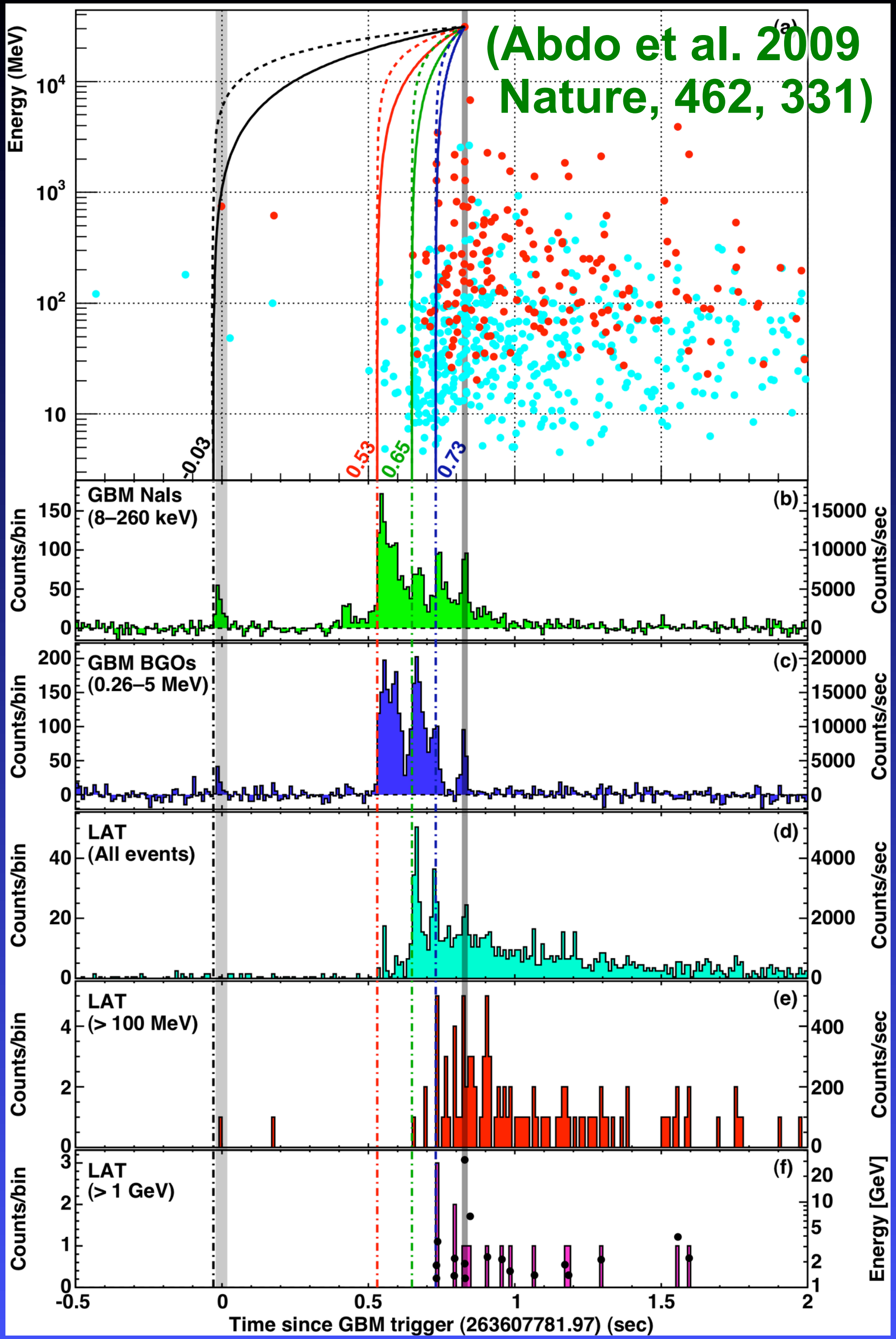


(Abdo et al. 2009, Science, 323, 1688)



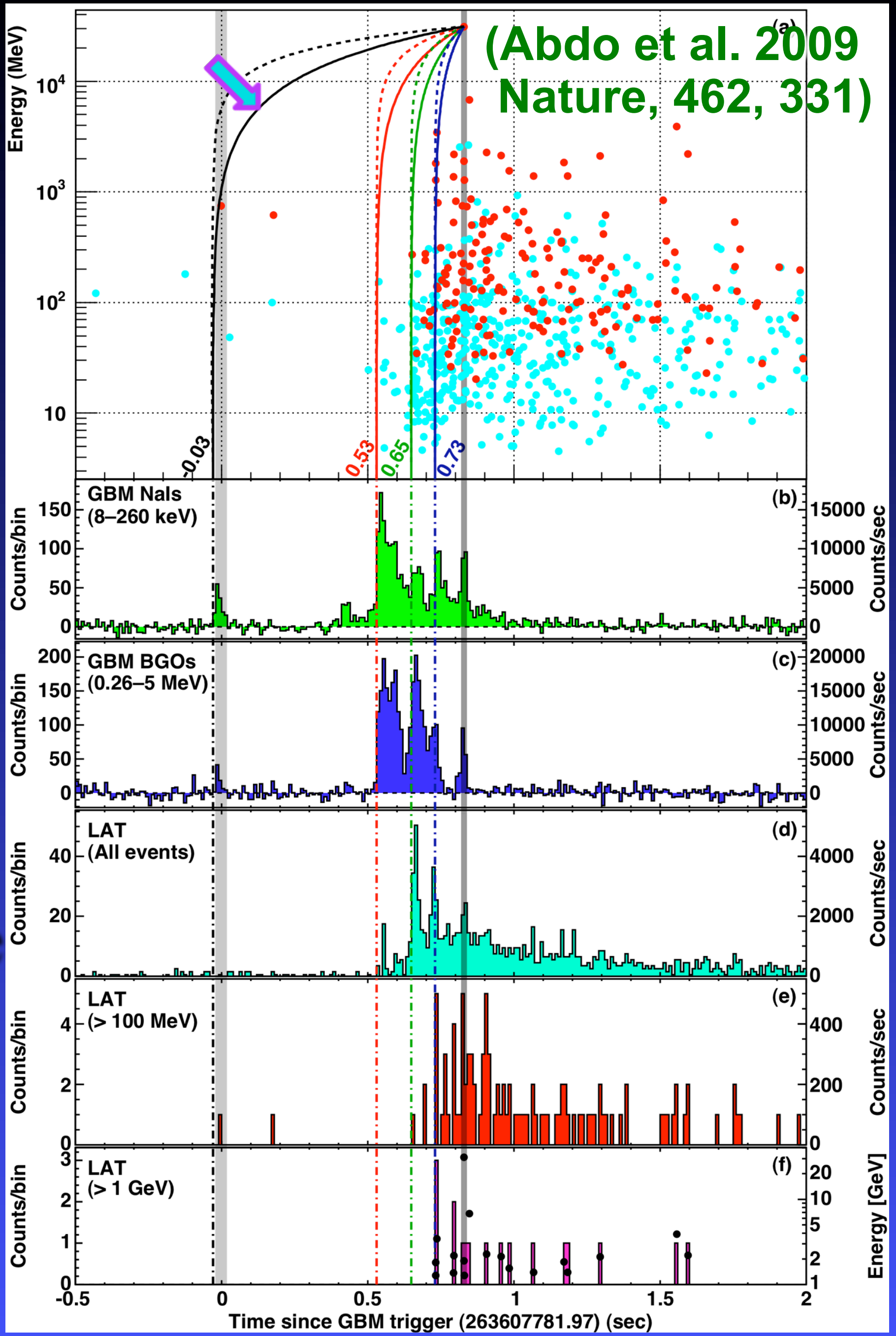
GRB090510: L.I.V

- A short GRB (duration ~1 s)
- Redshift: $z = 0.903 \pm 0.003$
- A ~31 GeV photon arrived at $t_h = 0.829$ s after the trigger
- We carefully verified it is a photon; from the GRB at $>5\sigma$
- We use the 1-σ lower bounds on the measured values of E_h (28 GeV) and z (0.900)
- Intrinsic spectral lags known on timescale of individual pulses: weak effect expected



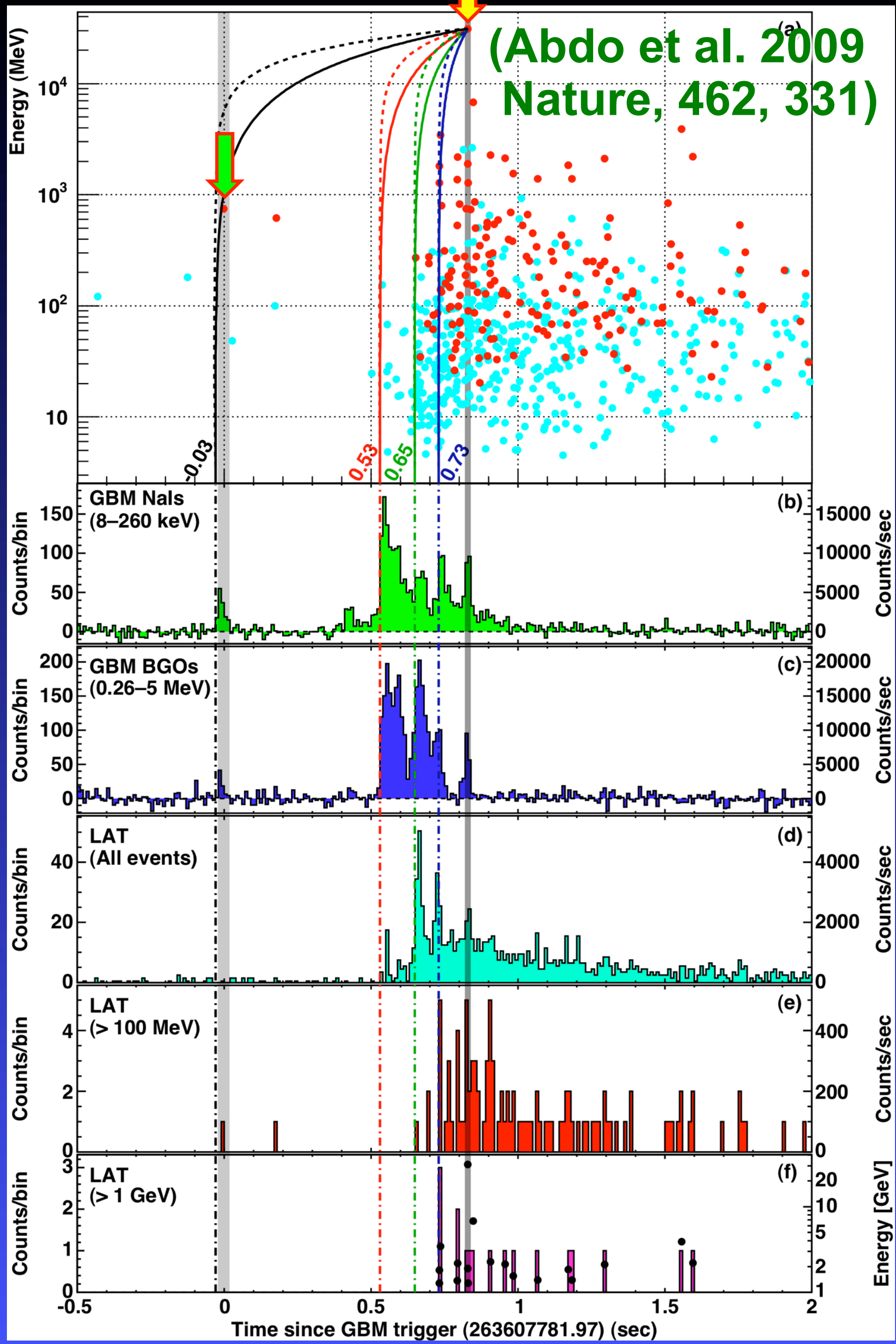
GRB090510: L.I.V

- Method 1: different choices of t_{start} from the most conservative to the least conservative
- $t_{start} = -0.03$ s precursor onset $\Rightarrow$ $\xi_1 = M_{QG,1}/M_{Planck} > 1.19$
- $t_{start} = 0.53$ s onset of main emission episode $\Rightarrow \xi_1 > 3.42$
- For any reasonable emission spectrum a ~31 GeV photon is accompanied by many γ's above 0.1 or 1 GeV that "mark" its t_{em}
- $t_{start} = 0.63$ s, 0.73 s onset of emission above 0.1, 1 GeV $\Rightarrow \xi_1 > 5.12$, $\xi_1 > 10.0$



GRB090510: L.I.V

- Method 2: **least conservative**
- Associating a high energy photon with a sharp spike in the low energy lightcurve, which it falls on top of
- Limits both signs: $s_n = \pm 1$
- Non-negligible chance probability (~5-10%), but still provides useful information
- For the 31 GeV photon (*shaded vertical region*) $\Rightarrow |\Delta t| < 10$ ms and $\xi_1 = M_{QG,1}/M_{Planck} > 102$
- For a 0.75 GeV photon during precursor: $|\Delta t| < 19$ ms, $\xi_1 > 1.33$



(Abdo et al. 2009 Nature, 462, 331)

Method 3: DisCan (Scargle et al. 2008)

- Based on lack of smearing of the fine time structure (sharp narrow spikes in the lightcurve) due to energy dispersion
- Constrains both possible signs of the effect: $s_n = \pm 1$
- Uses all LAT photons during the brightest emission episode (obs. range 35 MeV – 31 GeV); no binning in time or energy
- Shifts the arrival time of photons according to a trail energy dispersion (linear in our case), finding the coefficient that maximizes a measure of the resulting lightcurve variability
- We found a symmetric upper limit on a linear dispersion: $|\Delta t/\Delta E| < 30$ ms/GeV (99% CL) $\Rightarrow$ $M_{QG,1} > 1.22 M_{Planck}$
- Remains unchanged when using only photons < 3 or 1 GeV (a very robust limit)

Limits on LIV from Fermi GRBs

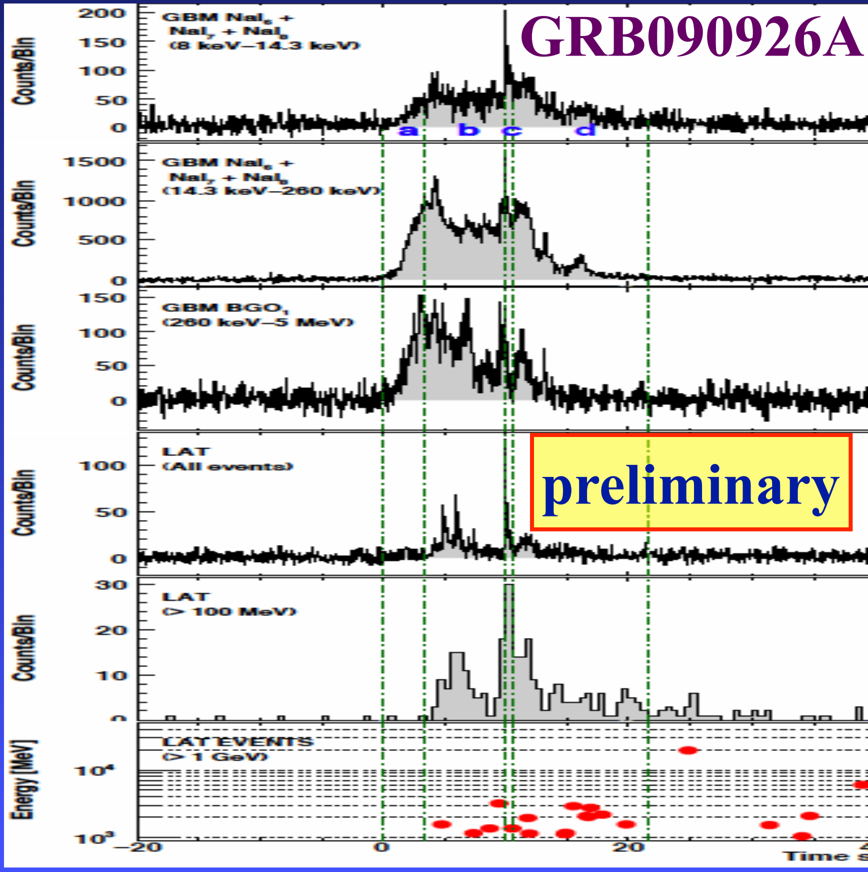
GRB	duration or class	# of events > 0.1 GeV	# of events > 1 GeV	method	Lower Limit on $M_{QG,1}$ / M_{Planck}	Valid for S_n =	Highest photon Energy	redshift
080916C	long	145	14	1	0.11	+1	~ 13 GeV	~ 4.35
090510	short	> 150	> 20	1 2 3	1.2, 3.4, 5.1, 10 102 1.2	+1 ±1 ±1	~ 31 GeV	0.903
090902B	long	> 200	> 30	1	0.068	+1	~ 33 GeV	1.822
090926	long	> 150	> 50	1 , 3	0.066, 0.082	+1	~ 20 GeV	2.1062

- Method 1: assuming a high-energy photon is not emitted before the onset of the relevant low-energy emission episode
- Method 2: associating a high-energy photon with a spike in the low-energy light-curve that it coincides with
- Method 3: DisCan (dispersion cancelation; very robust) – lack of smearing of narrow spikes in high-energy light-curve

Prospects for LIV studies with CTA GRBs

- Method 1: it may be difficult to do much better
  - Our current limit $|\Delta t/\Delta E| < 30$ ms/GeV would require $E_h > 1$ TeV for a response time of 30 s
  - at > 1 TeV intrinsically fewer photons + EBL
- Method 3: might work best
  - Sharp bright spikes up to high energies exist also well within long GRBs
  - $t_{var} \sim 0.1$ s & $E_h \sim 0.1$ TeV could do ~30 times better
- A short GRB in CTA FoV (survey mode) would be great

10 ms, 1 TeV: >10^3 times better



Conclusions:

- GRBs are very useful for constraining LIV
- Bright **short** GRBs are more useful than long ones
- A very robust and conservative limit on a linear energy dispersion of either sign: $M_{QG,1} > 1.22 M_{Planck}$
- Still conservative but somewhat less robust limits: $M_{QG,1}/M_{Planck} > 5.1,\ 10$ (onset of emission >0.1, 1 GeV)
- "Intuition builder": $M_{QG,1} / M_{planck} > 102$
- Quantum-Gravity Models with linear energy dispersion ($n = 1$) are disfavored