Lessons from GW170817 / GRB170817A

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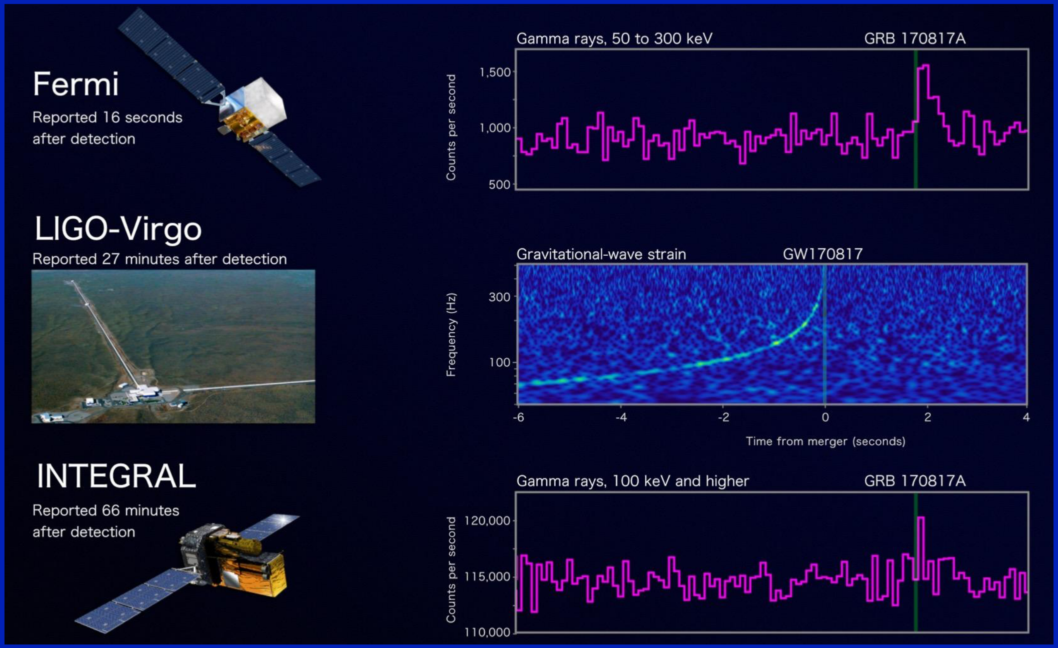

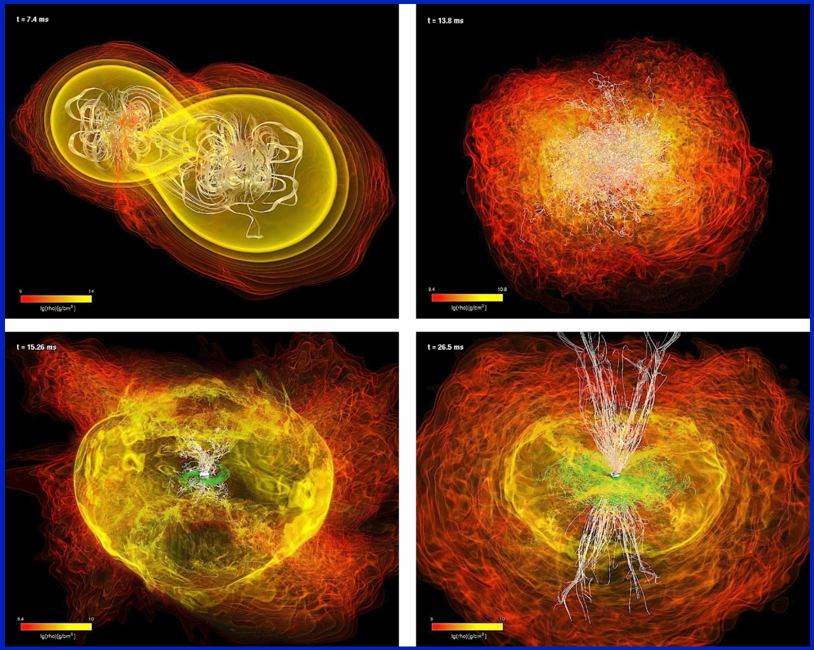

Forum for Multi-Wavelength Survey and Time-Domain Astronomy

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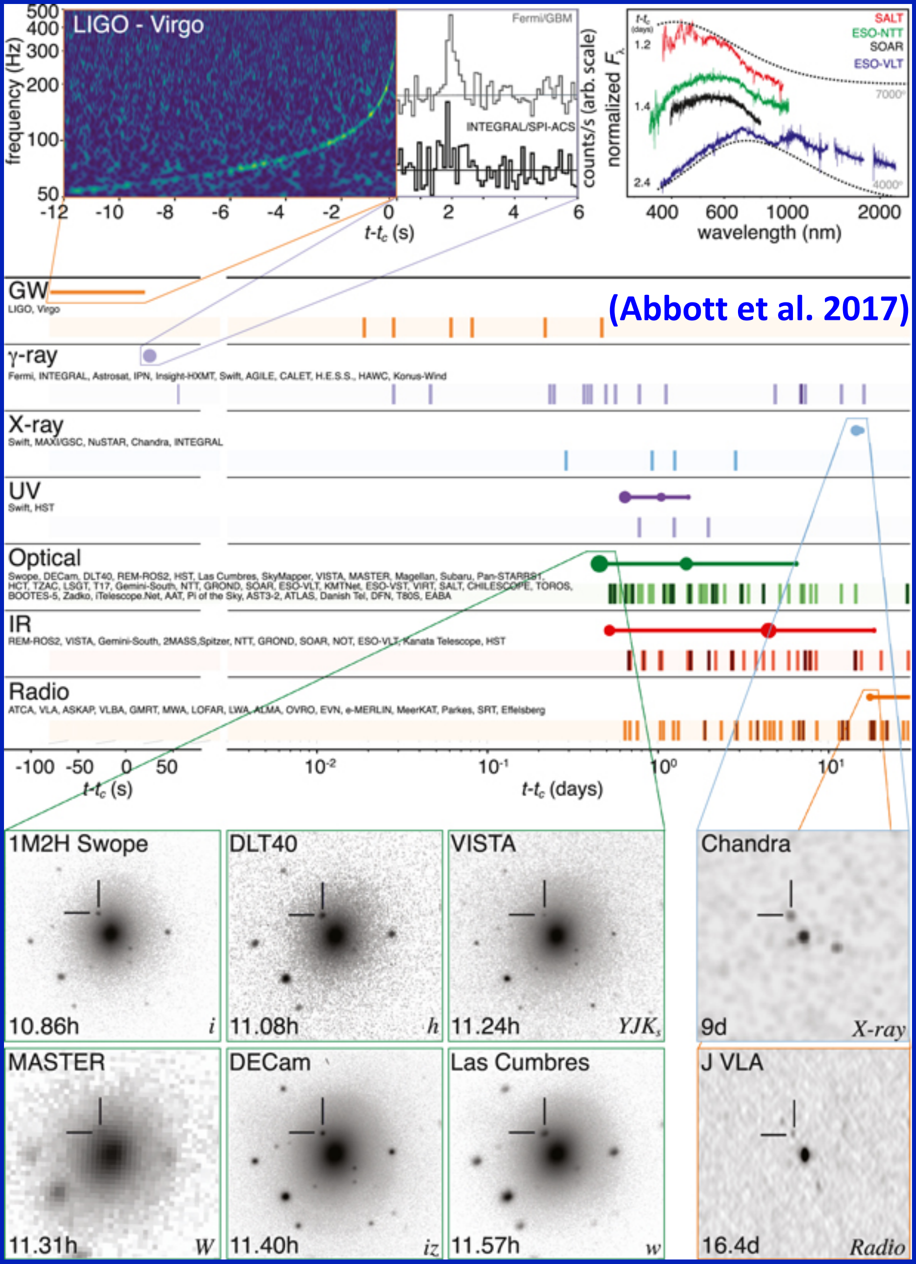
Outline of the Talk:

- The **Extraordinary** event **GW170817/GRB170817A**
- **The merger remnant: Black Hole or a massive NS?**
- The **afterglow** emission:
  - **Two main options for the early flux rise: r vs. θ dependence**
  - **Breaking the degeneracy: lightcurves? Images, Polarization**
  - **New observations imply: dominant θ dependence (off-axis jet)**
  - Radio polarization U.L. constrains the shock-produced B-field
- **Conclusions**

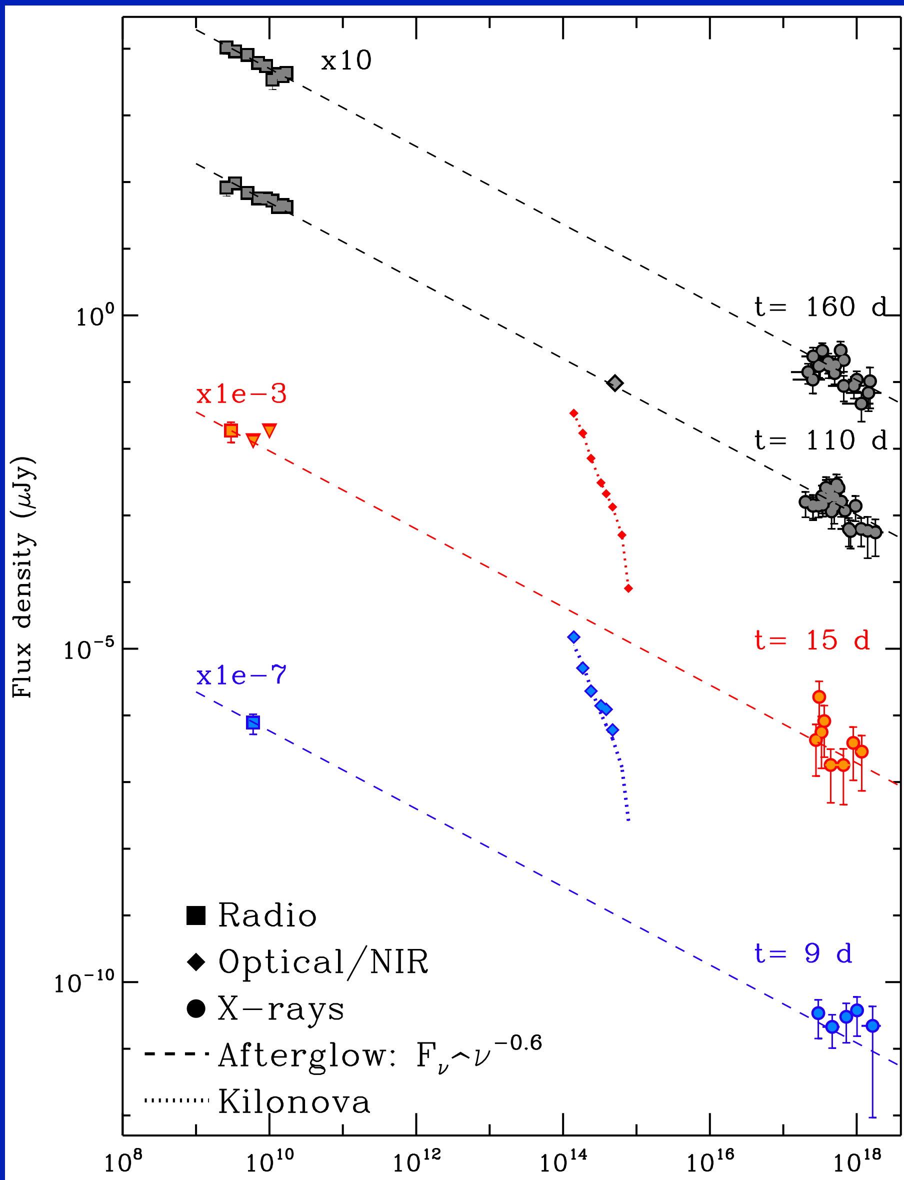
GW 170817 / GRB 170817A: D ≈ 40 Mpc

- **First GW detection** of a **NS-NS merger**
- **First electromagnetic counterpart** to a GW event
  - The short GRB 170817A (very under-luminous, 1.74 s γ-GW delay)
  - Optical (IR to UV) kilonova emission over a few weeks
  - X-ray (> 9 d) to radio (>16 d) afterglow (still detected)
- **First direct association of a sGRB & NS-NS merger***
  (Eichler+ 1989; Narayan+ 1992)
- **First clear-cut kilonova**

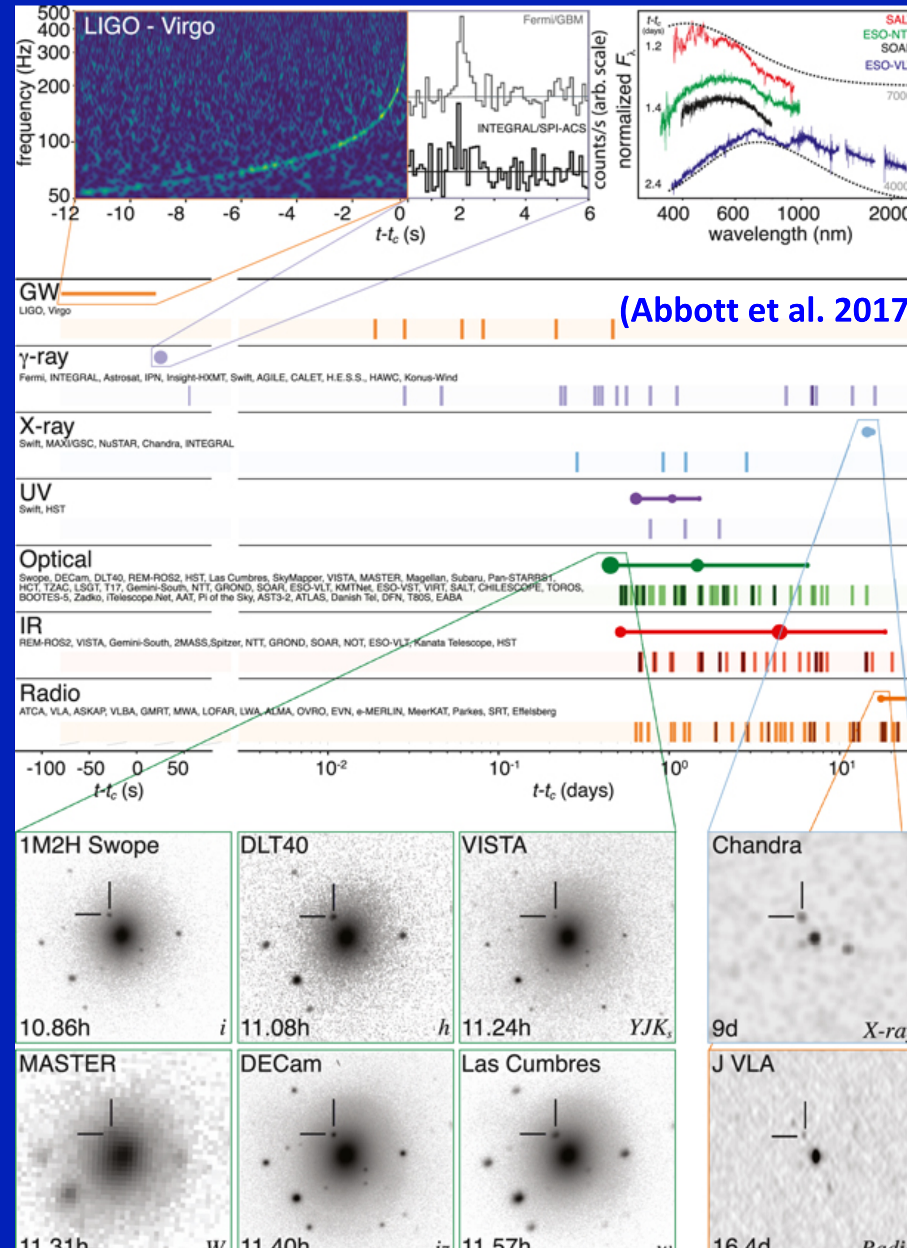


(Abbott et al. 2017)

GW 170817 / GRB 170817A: D ≈ 40 Mpc

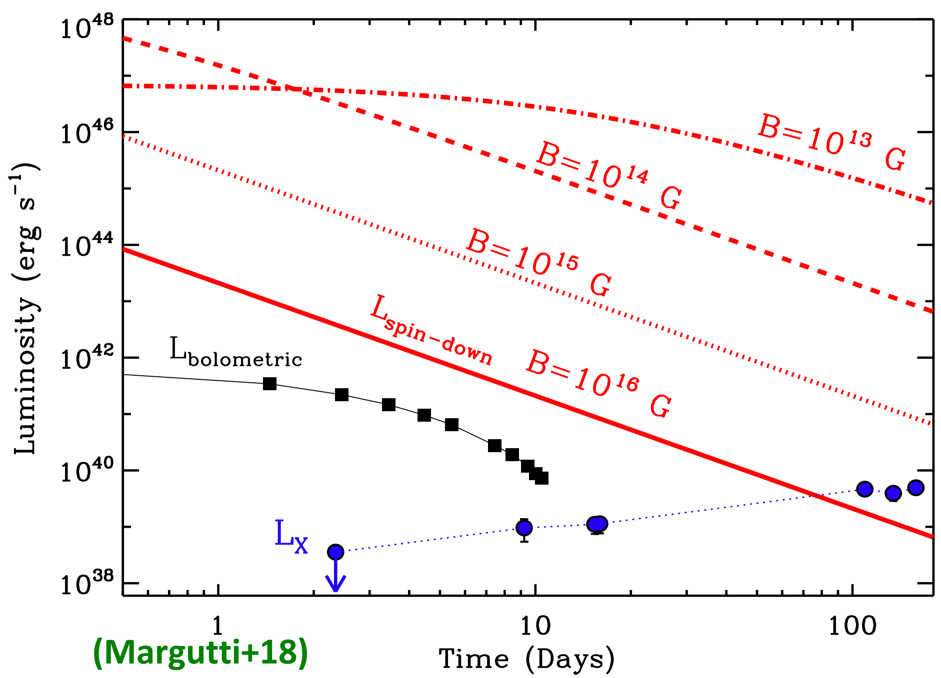


(Margutti et al. 2018)



GW 170817: the type of **remnant**

- $M_{1,2}$ = pre-merger NS $M_{gravitational}$
post-merger total mass: $M_i = M_1 + M_2$
- Final mass $M_f \approx 0.93 M_i$ due to:
  - GW & neutrino energy losses
  - Mass ejection during the merger
- A stable NS or SMNS $\Rightarrow P_0 \approx 1$ ms
$\Rightarrow E_{rot} \gtrsim 10^{52.5}$ erg, $\tau_{sd} \approx 20 B_{13}^{-2}$ days
$\Rightarrow$ would contradict afterglow obs.
(also what produces the GRB/afterglow?)

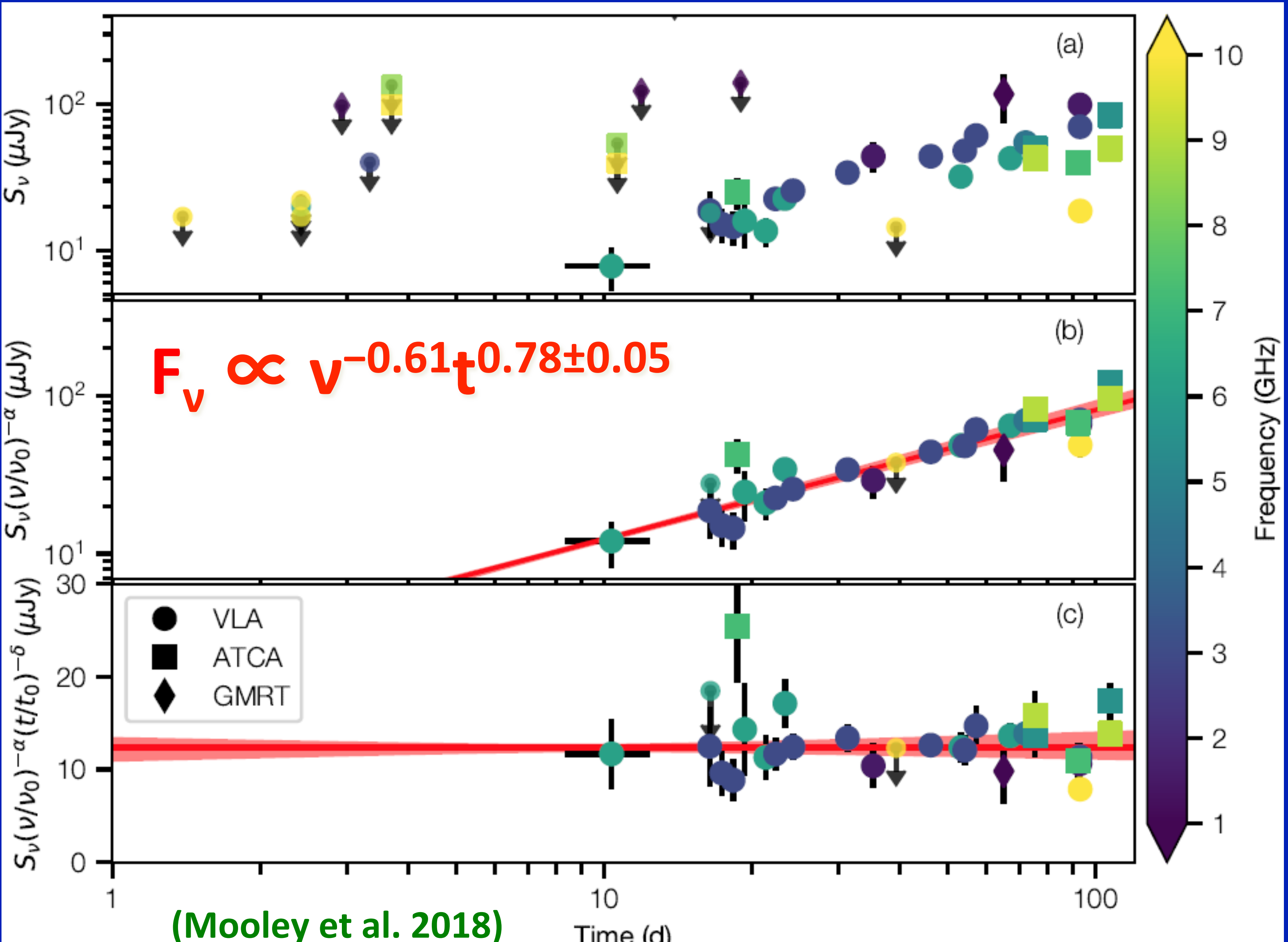


(Margutti+18)

Chirp mass: $\mathcal{M} = \left(\frac{M_1^3 M_2^3}{M_1 + M_2}\right)^{1/5} = 1.188^{+0.004}_{-0.002} M_\odot$ (Abbott+ 2017)

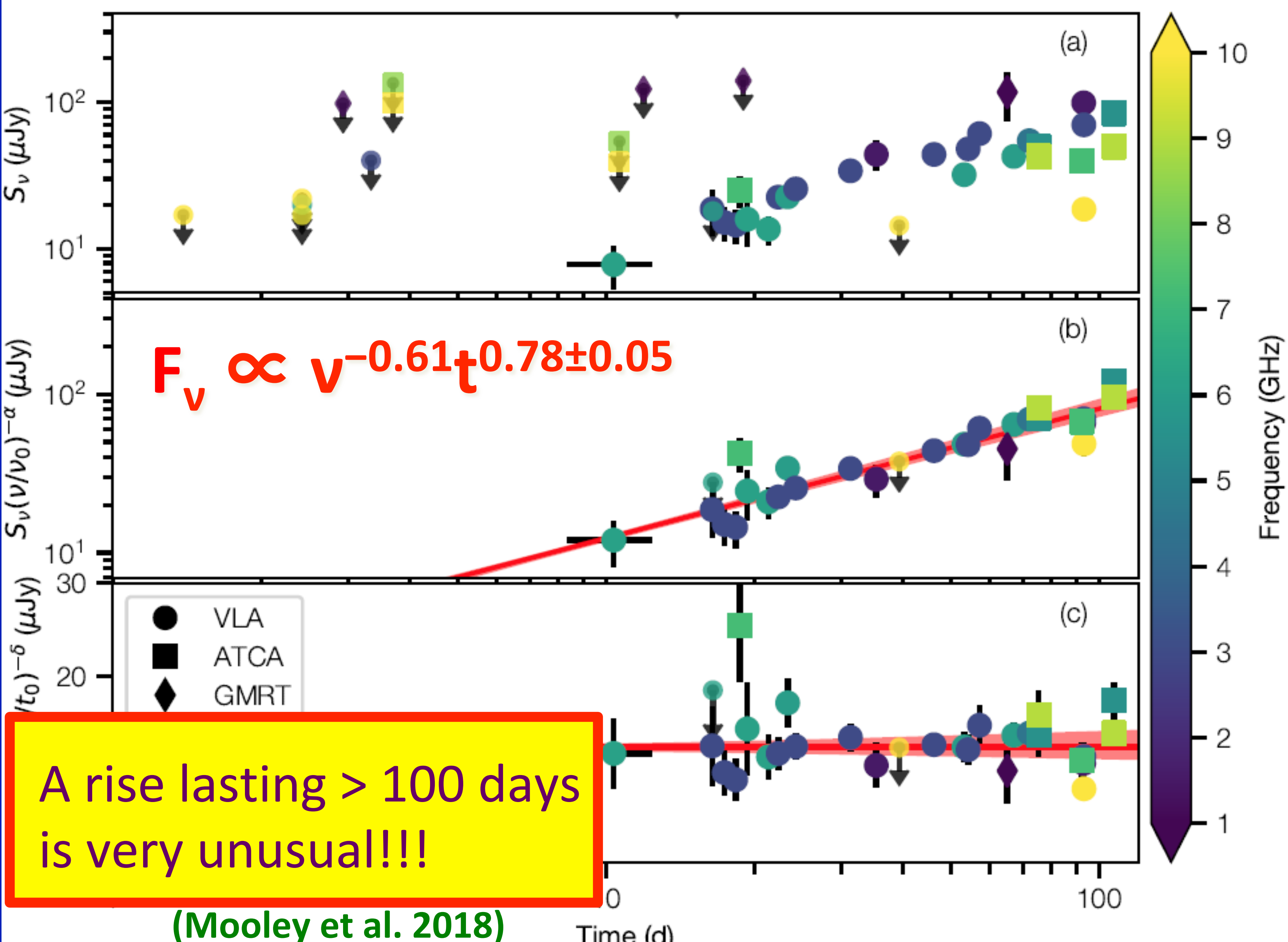
- The argument can be reversed to constrain NS EoS & $M_{max} \lesssim 2.17 M_\odot$
(Margalit & Metzger 2017; Rezzolla et al. 2018)

GRB 170817A: afterglow observations



(Mooley et al. 2018)

GRB 170817A: afterglow observations



$F_\nu \propto \nu^{-0.61} t^{0.78\pm0.05}$

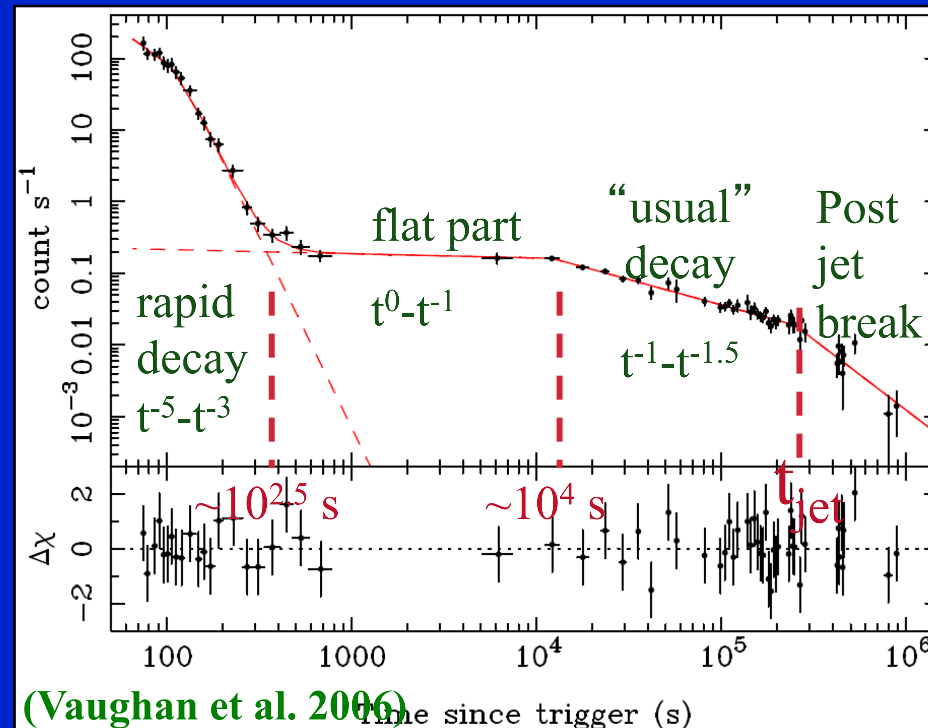
A rise lasting > 100 days is very unusual!!!

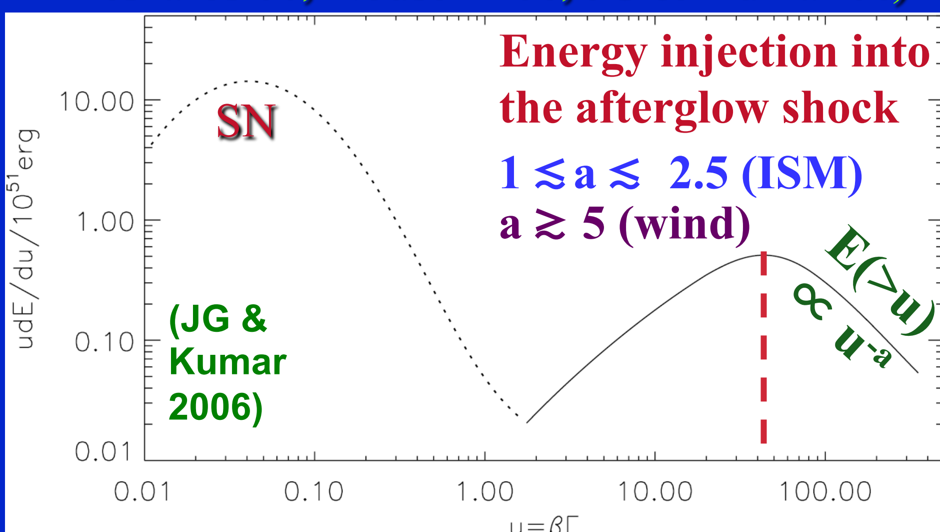
(Mooley et al. 2018)

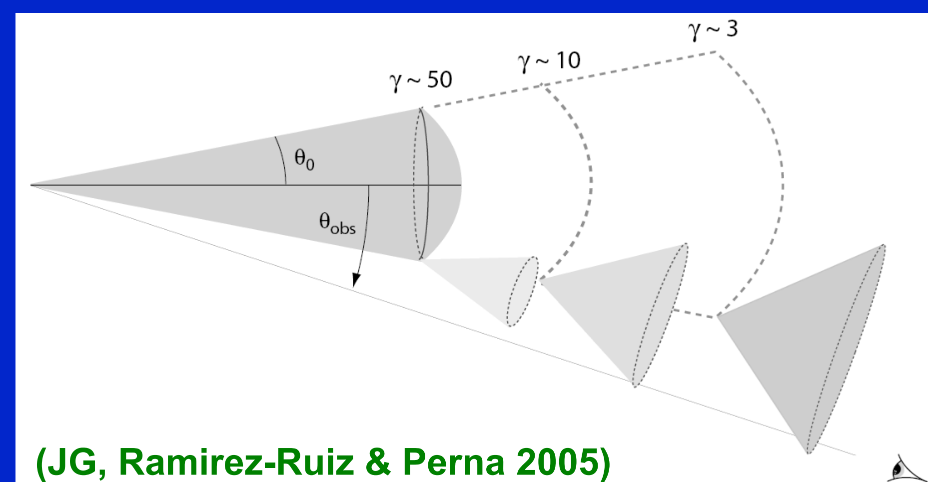
Analogy to rising F_ν: X-ray Plateaus

- Possible solutions:
- Evolution of shock microphysical parameters (JG, Konigl & Piran 2006)
- Energy injection into ext. shock:

1. long-lived relativistic wind
2. slower ejecta catching up (Sari & Meszaros 00; Nousek+ 06; JG & Kumar 06)







- Viewing angle effects

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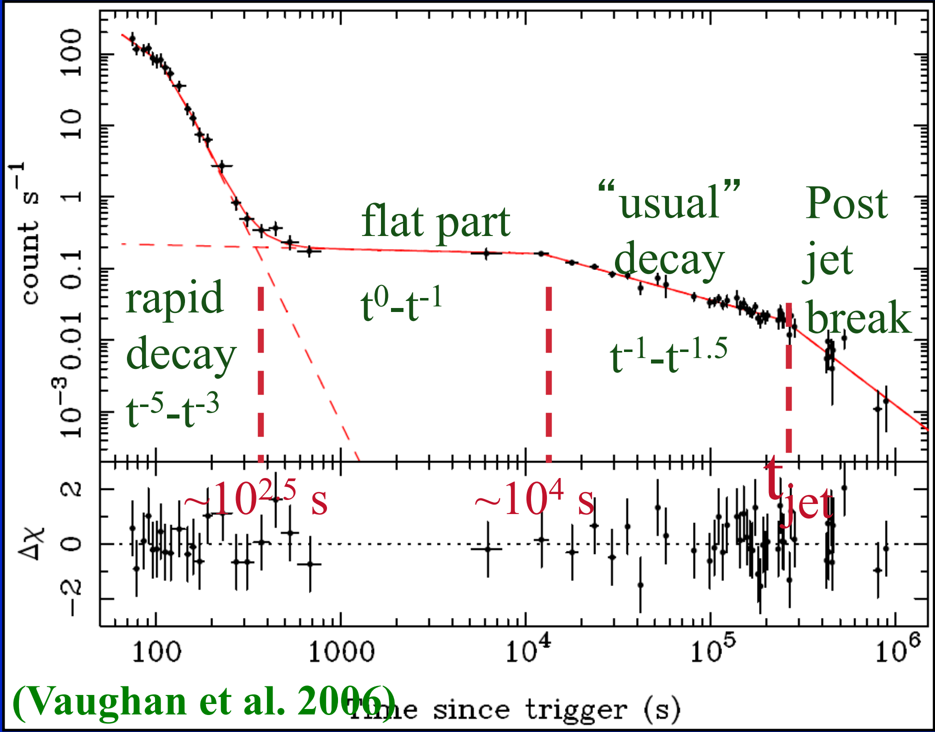
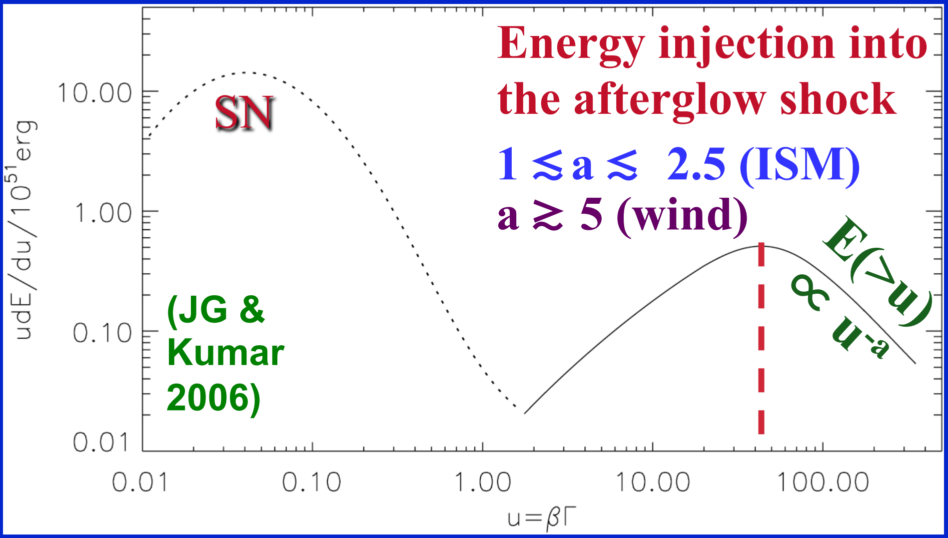
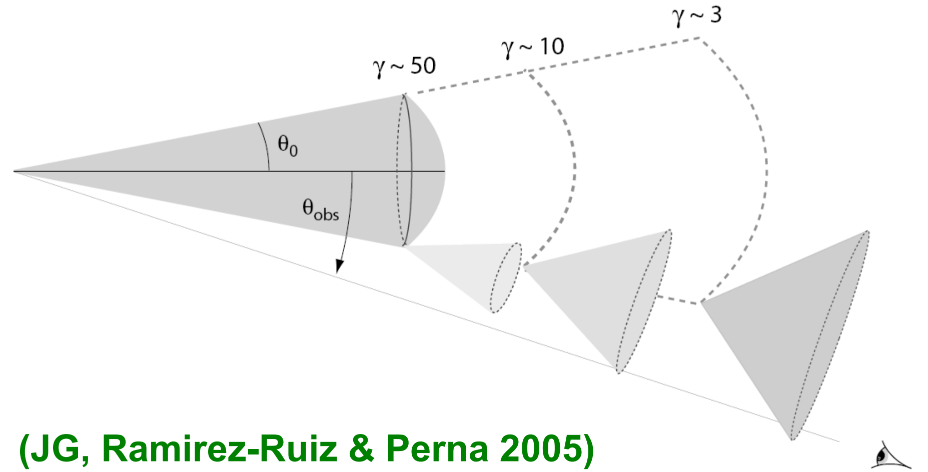


(Vaughan et al. 2006)



(JG & Kumar 2006)



(JG, Ramirez-Ruiz & Perna 2005)

- Viewing angle effects

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- Evolution of shock microphysical parameters (JG, Konigl & Piran 2006)
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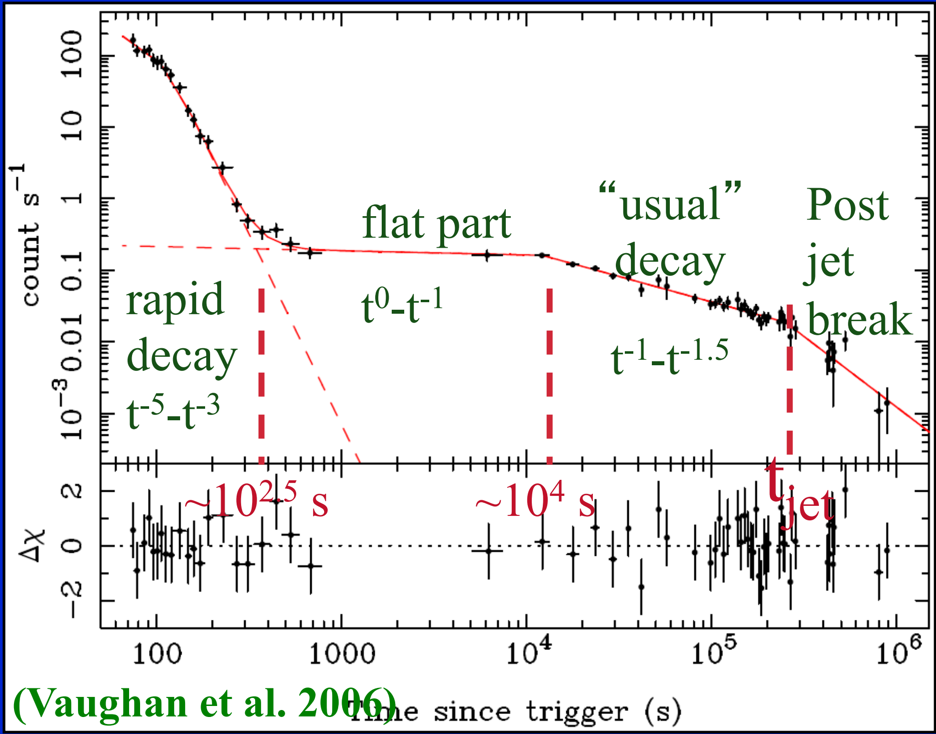
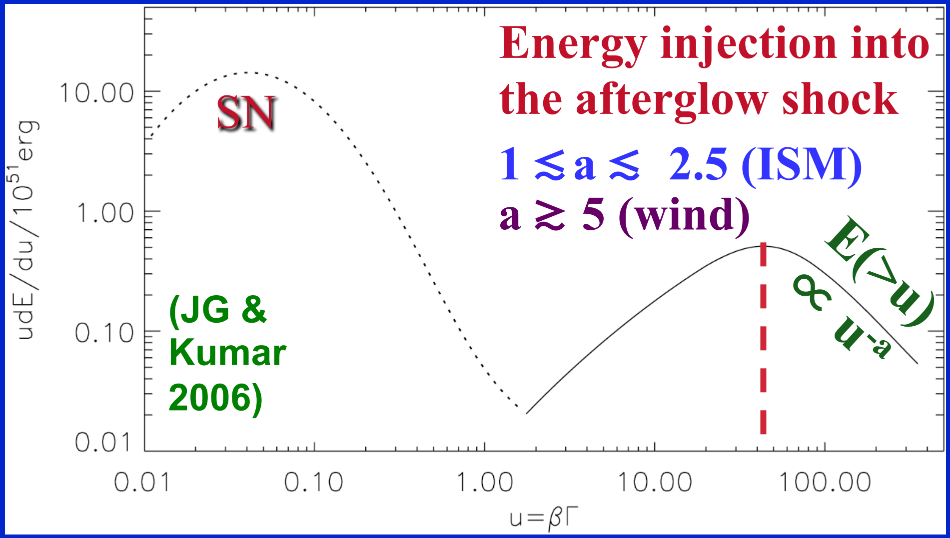
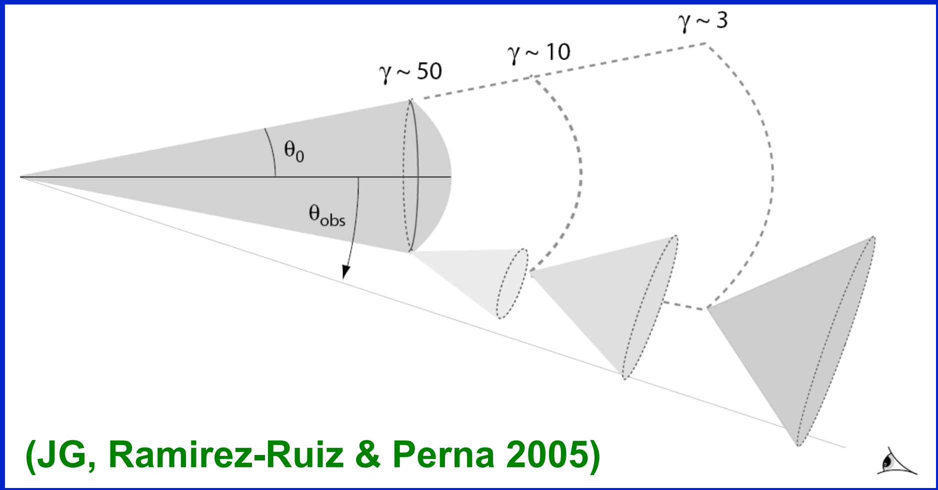




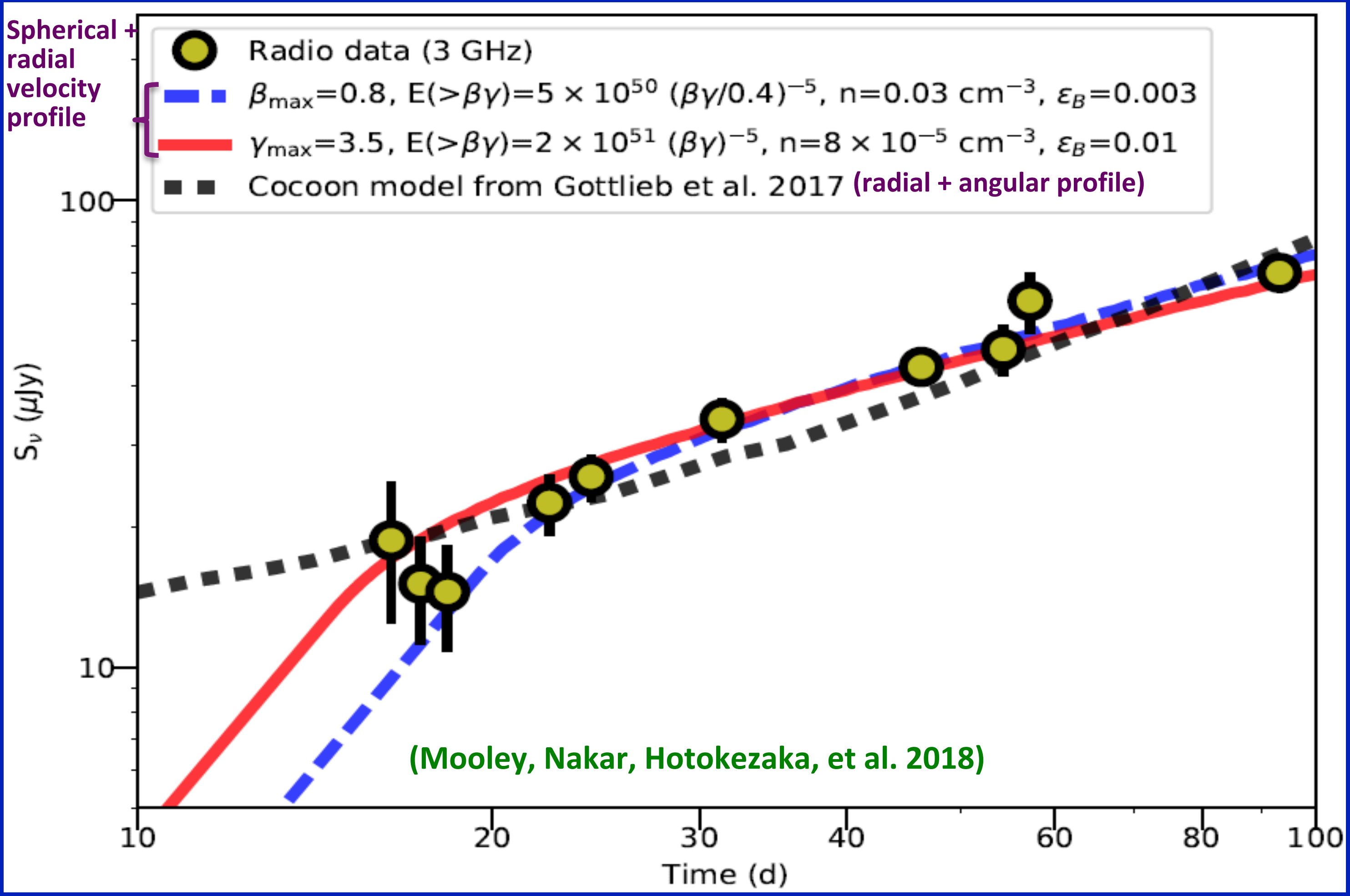


- Viewing angle effects **angular**

GRB170817 outflow structure: prompt, afterglow

- Cocoon model (Kasliwal+17; Mooley+18; Nakar & Piran 18): **r** & **θ** profile

◆ Cocoon-driven shock breakout can naturally produce the γ-rays (Kasliwal+17; Gottlieb+17; Bromberg+18; Nakar & Piran 18; Nakar+18)

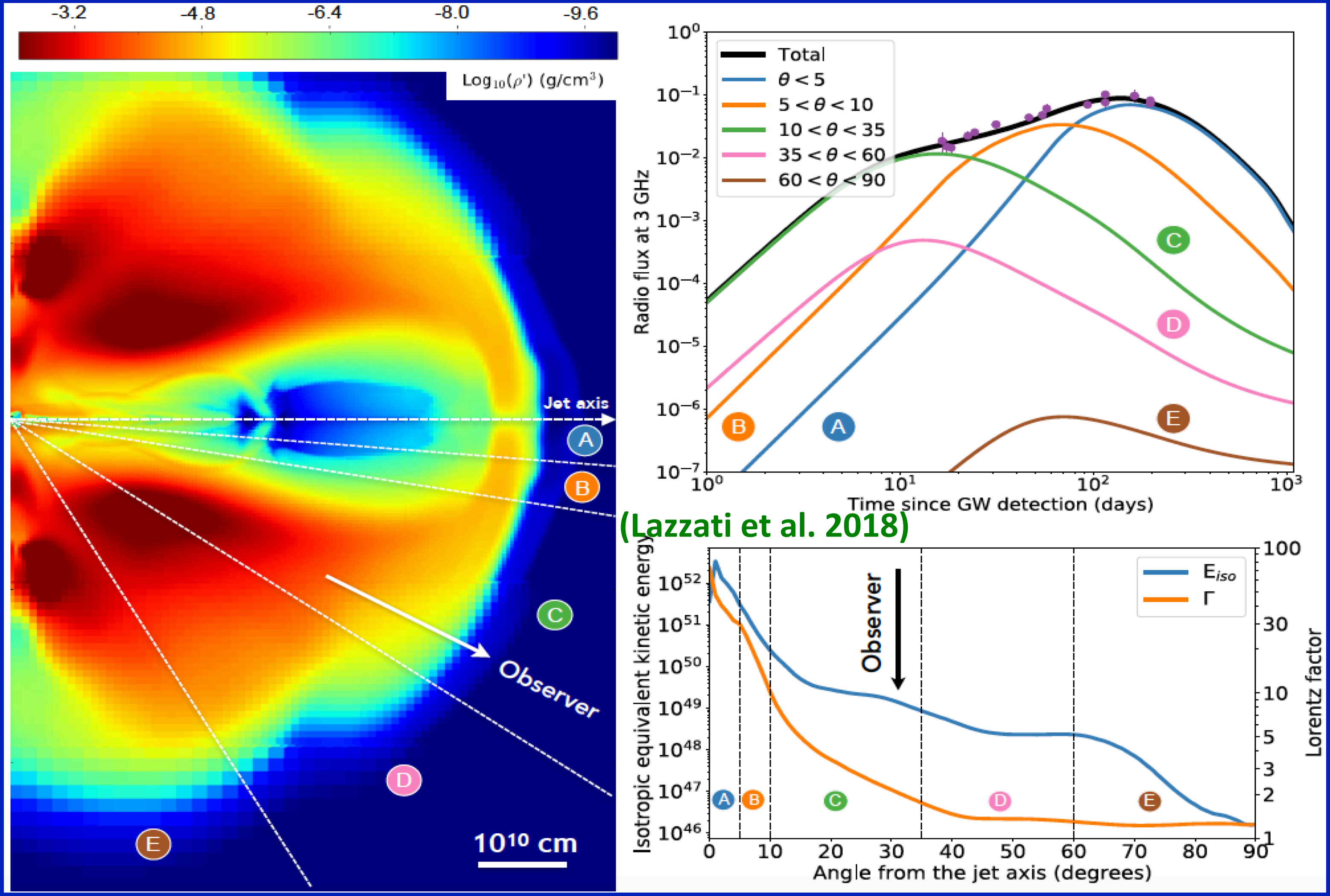


GRB170817 outflow structure: the afterglow

- A structured jet explanation (Lazzati+17; Margutti+18; Gill & JG 18;…):

◆ Simulation of jet breaking out of the Newtonian ejecta near a NS-NS merger site: the cocoon energizes the jet's sides/wings

◆ Afterglow dominated by θ profile

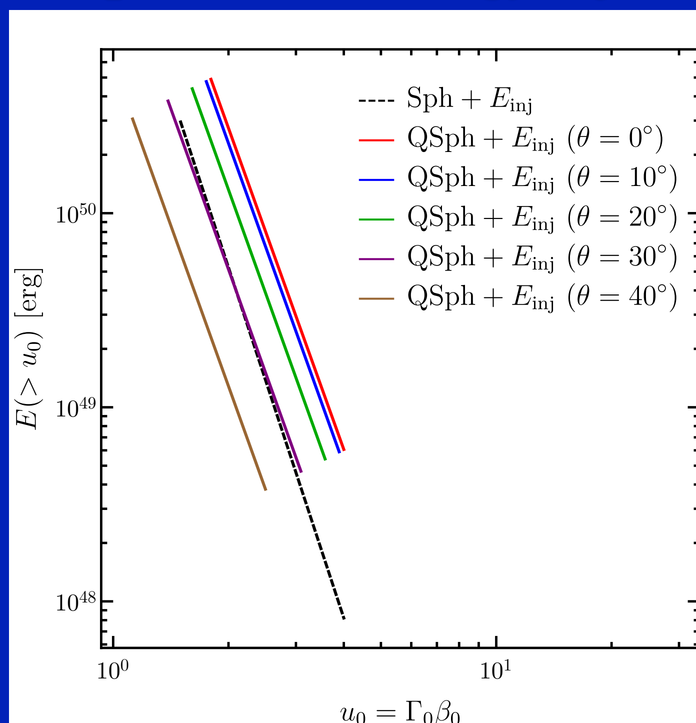


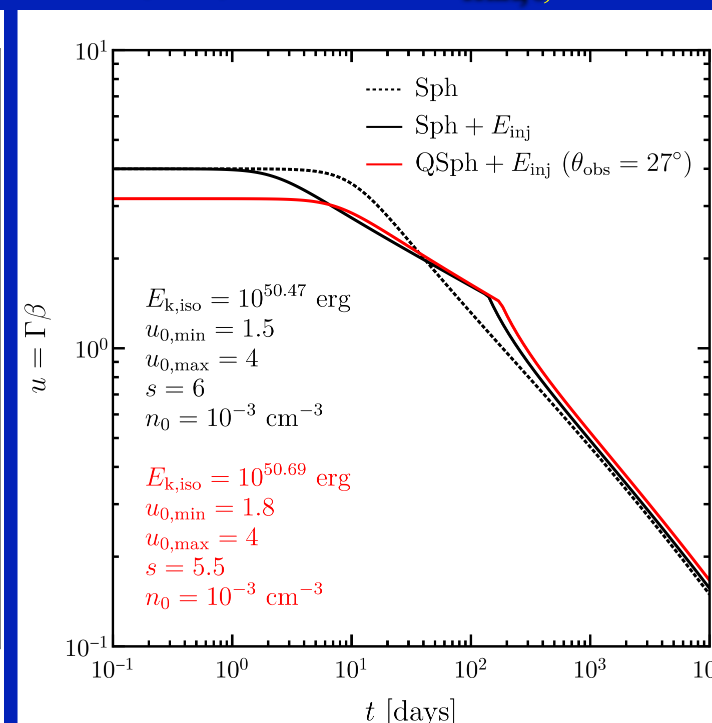
(Lazzati et al. 2018)

Outflow structure: breaking the degeneracy (Gill & JG 18)

- The lightcurves leave a lot of degeneracy between models
- The degeneracy may be lifted by calculation the afterglow images & polarization (e.g. Nakar & Piran 2018; Nakar et al. 2018)
- We considered 4 different models including both main types
  - Sph+E_{inj}: Spherical with energy injection $E(>u=\Gamma\beta) \propto u^{-6}$, $1.5<u<4$
  - QSph+E_{inj}: Quasi-Spherical+energy injection $E(>u) \propto u^{-s}$, $u_{min,0}=1.8$ $u_{max,0}=4$, $s = 5.5$, $\zeta = 0.1$

$$\frac{\epsilon(\theta)}{\epsilon_0} = \frac{u_{0,\min}(\theta)}{u_{\min,0}} = \frac{u_{0,\max}(\theta)}{u_{\max,0}} = \frac{\zeta + \cos^2\theta}{\zeta + 1}$$





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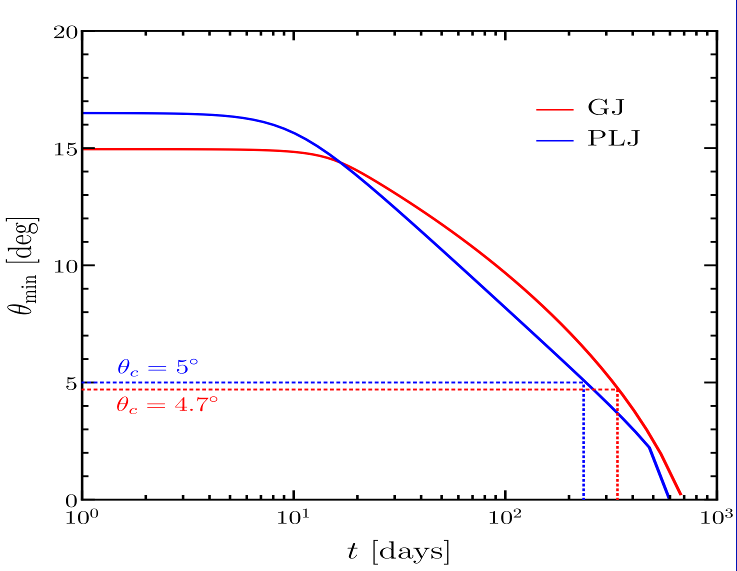
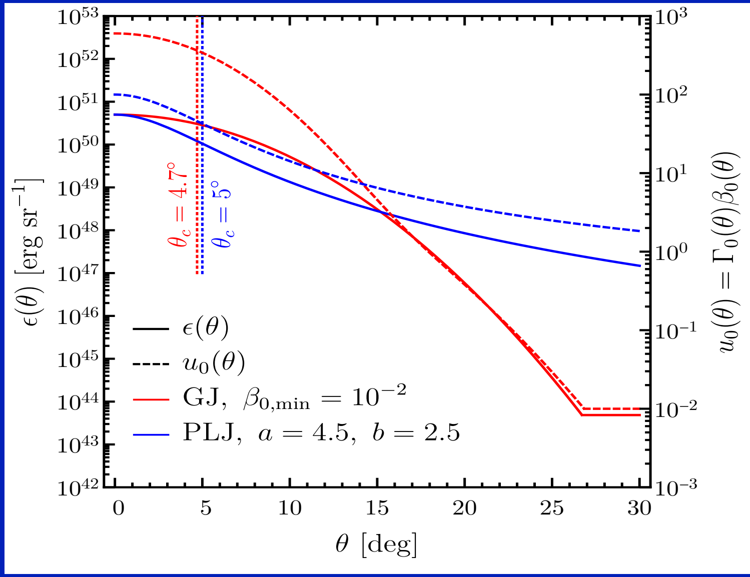
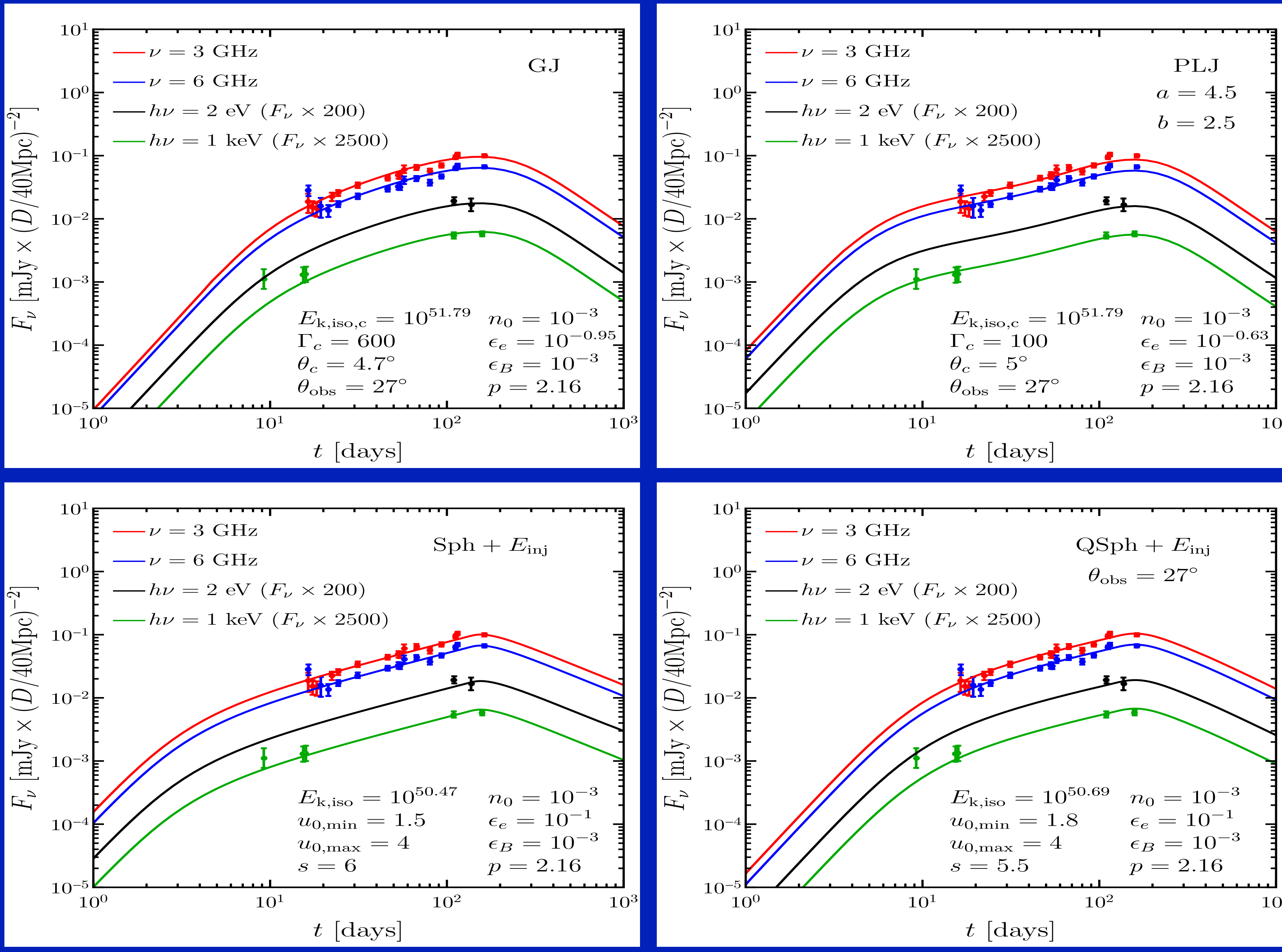
- The degeneracy may be lifted by calculation the afterglow images & polarization (e.g. Nakar & Piran 2018; Nakar et al. 2018)
- We considered 4 different models including both main types
  - GJ: Gaussian Jet (in $\varepsilon = dE/d\Omega$, Γ_0-1) $\Gamma_c = 600$, $\theta_c = 4.7°$
  - PLJ: Power-Law Jet; $\varepsilon = \varepsilon_c\Theta^{-a}$, $\Gamma_0-1 = (\Gamma_c-1)\Theta^{-b}$, $\Theta = [1+(\theta/\theta_c)^2]^{1/2}$ $\Gamma_c = 100$, $\theta_c = 5°$, $a = 4.5$, $b = 2.5$
- As there is a lot of freedom we fixed: $p = 2.16$, $\varepsilon_B = n_0 = 10^{-3}$, $\theta_{obs} = 27°$





The outflow structure: breaking the degeneracy

- Tentative fit to GRB170817A afterglow data (radio to X-ray)



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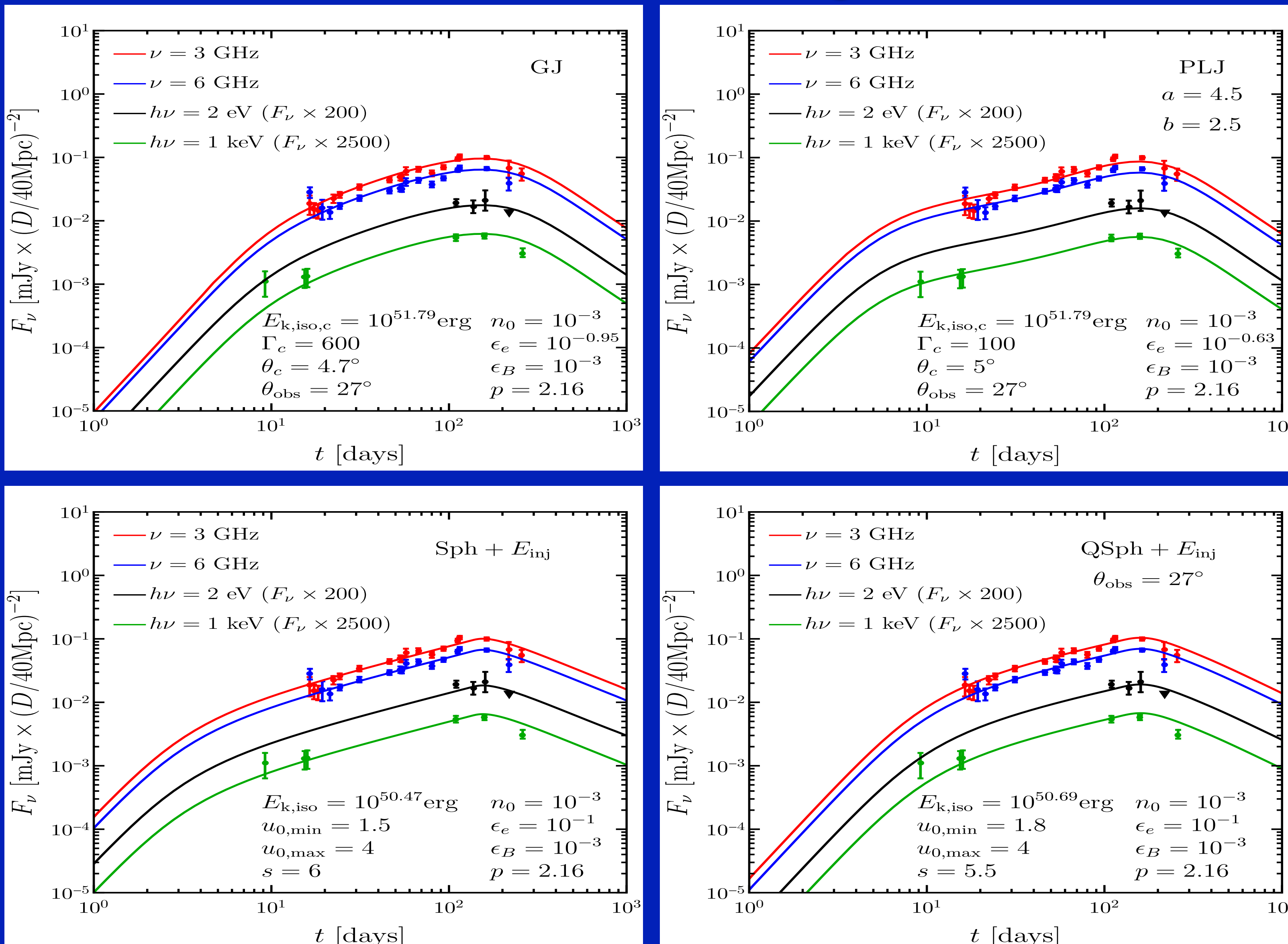
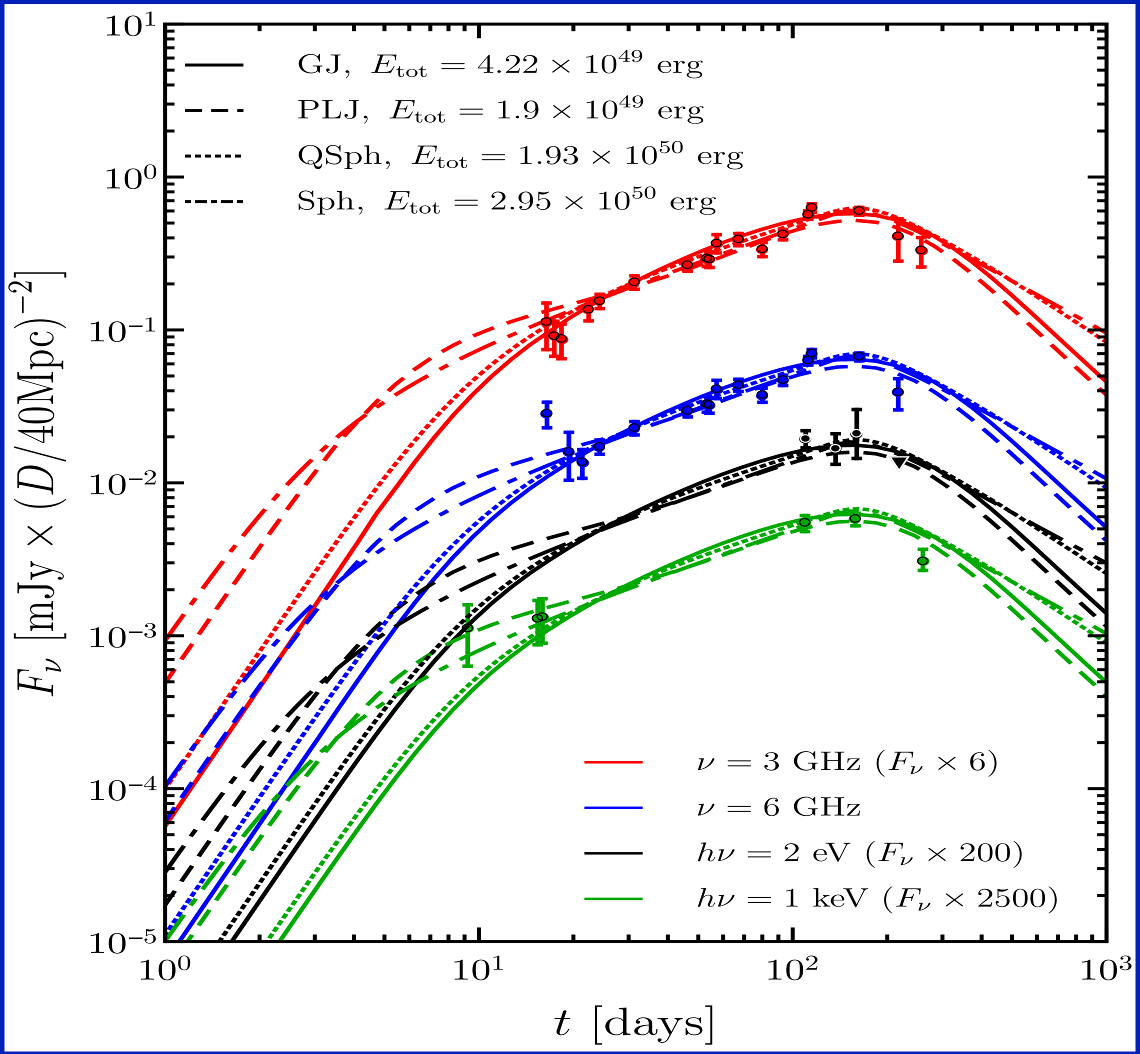


The outflow structure: breaking the degeneracy

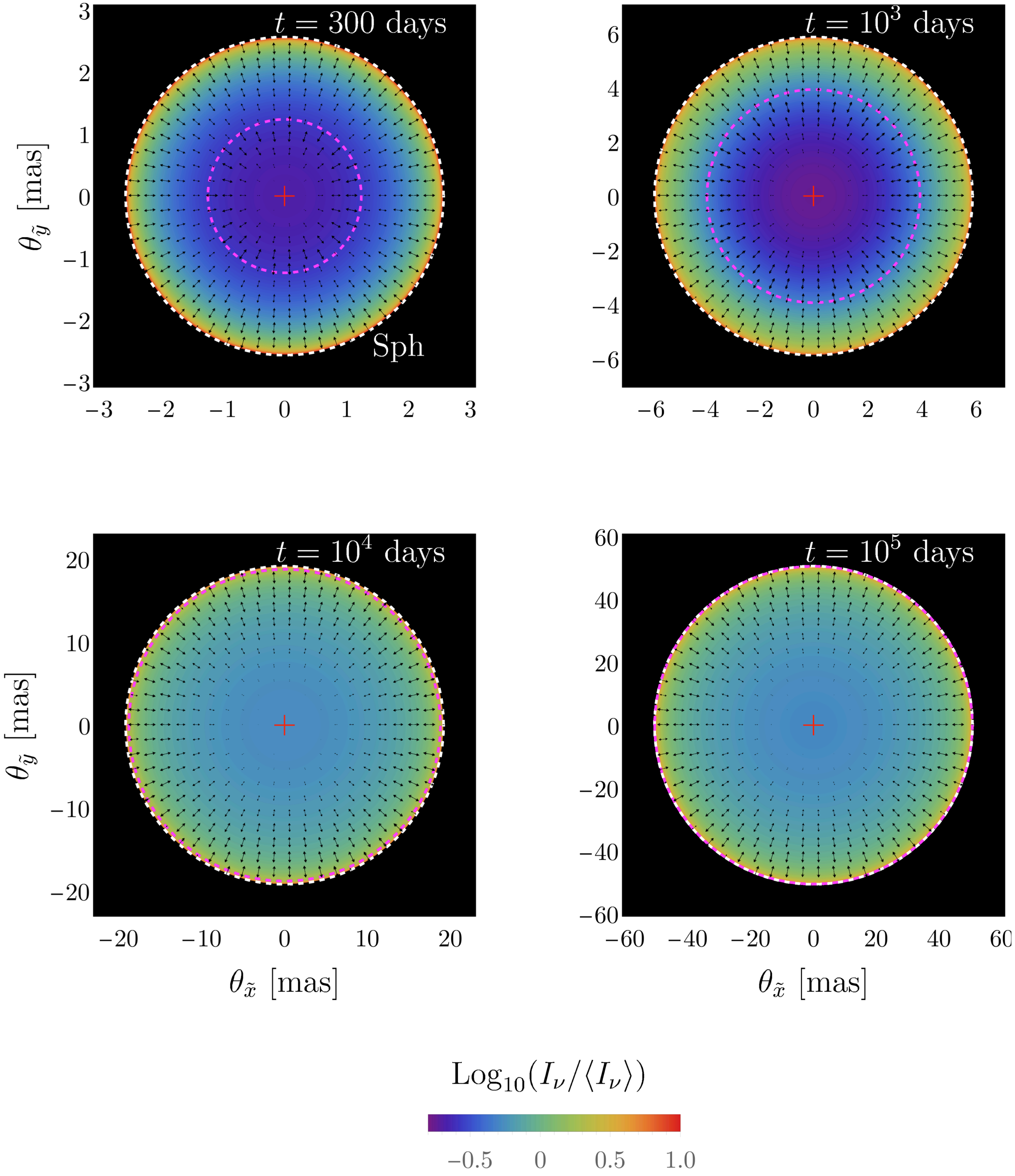
- New data that came out established a peak at t_p ~150 days
- The jet models decay faster (slightly preferred by the latest data)



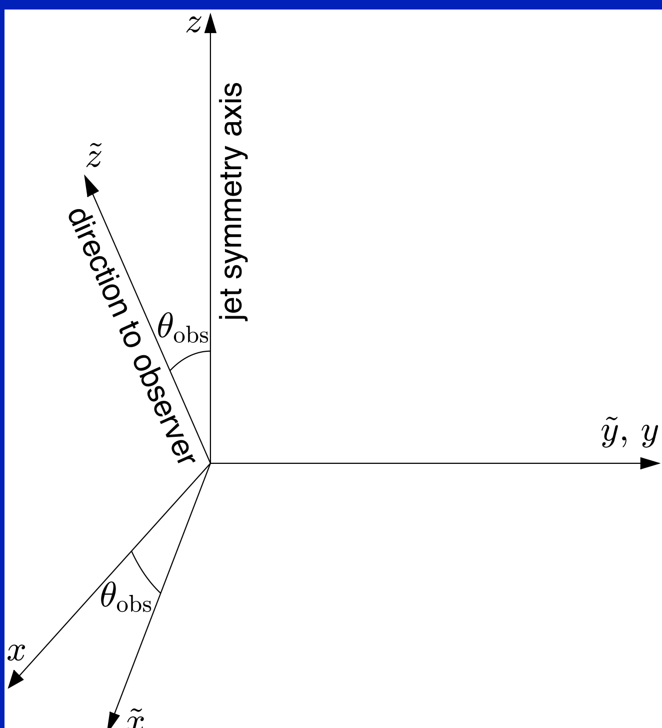
Decay: $\sim t^{-2.2}$

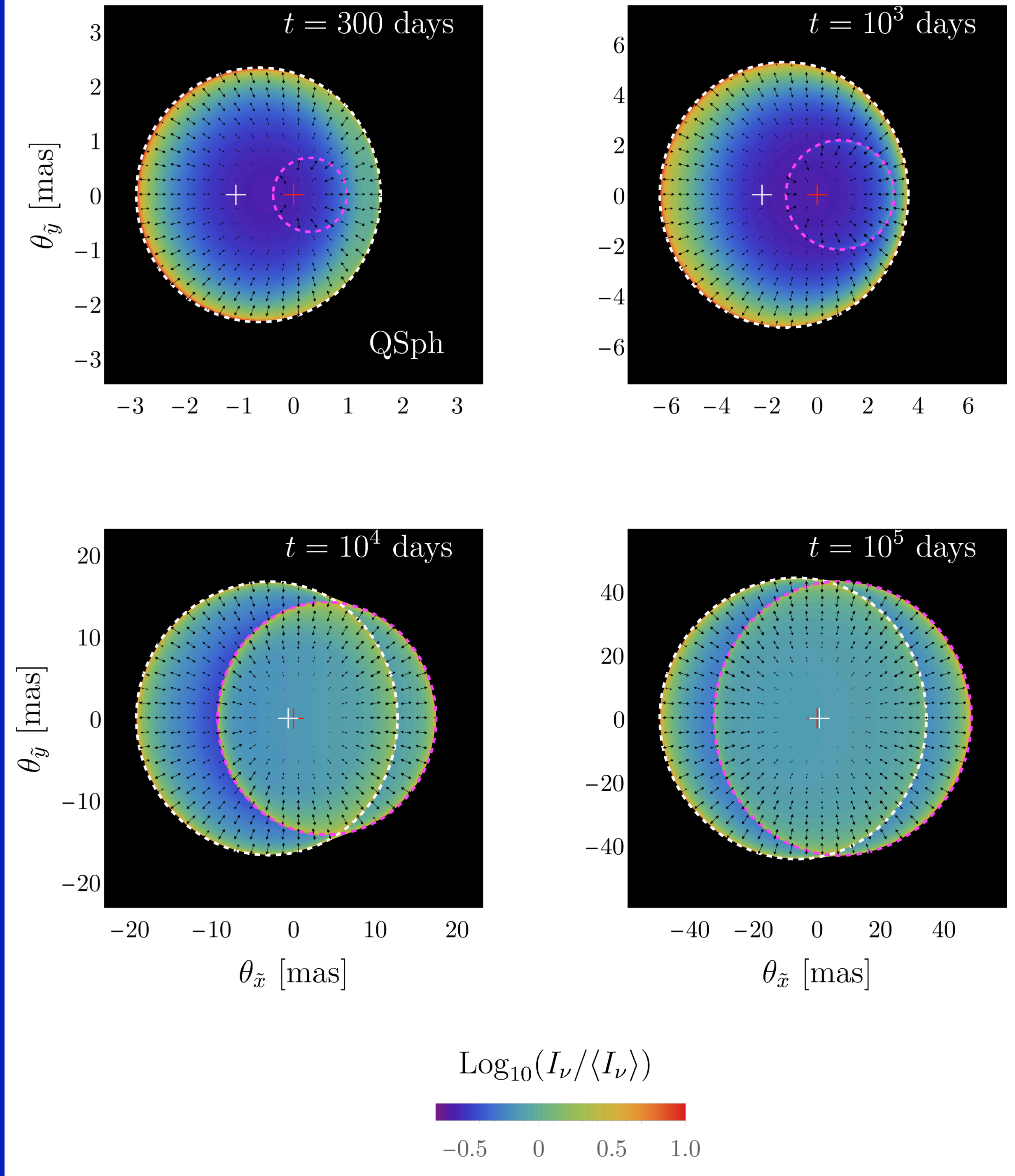
(Mooley+18)

Afterglow Images: Sph + E_{inj}

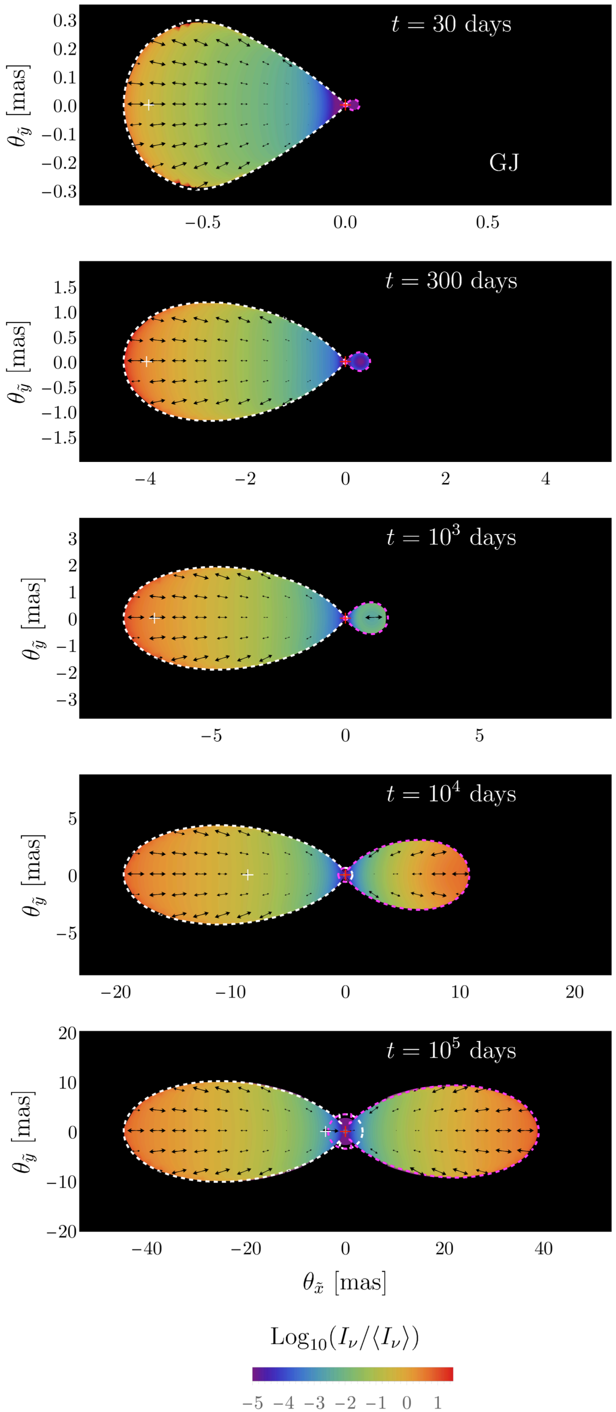


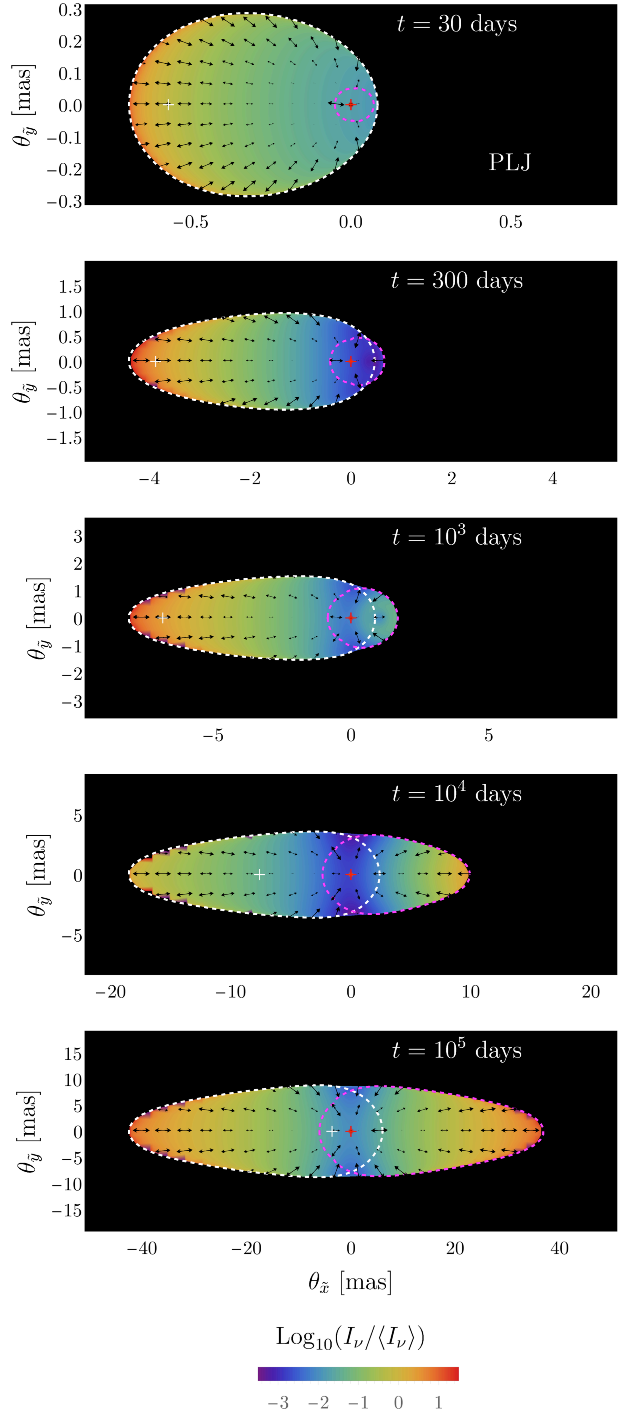
Afterglow Images: QSph + E_{inj}





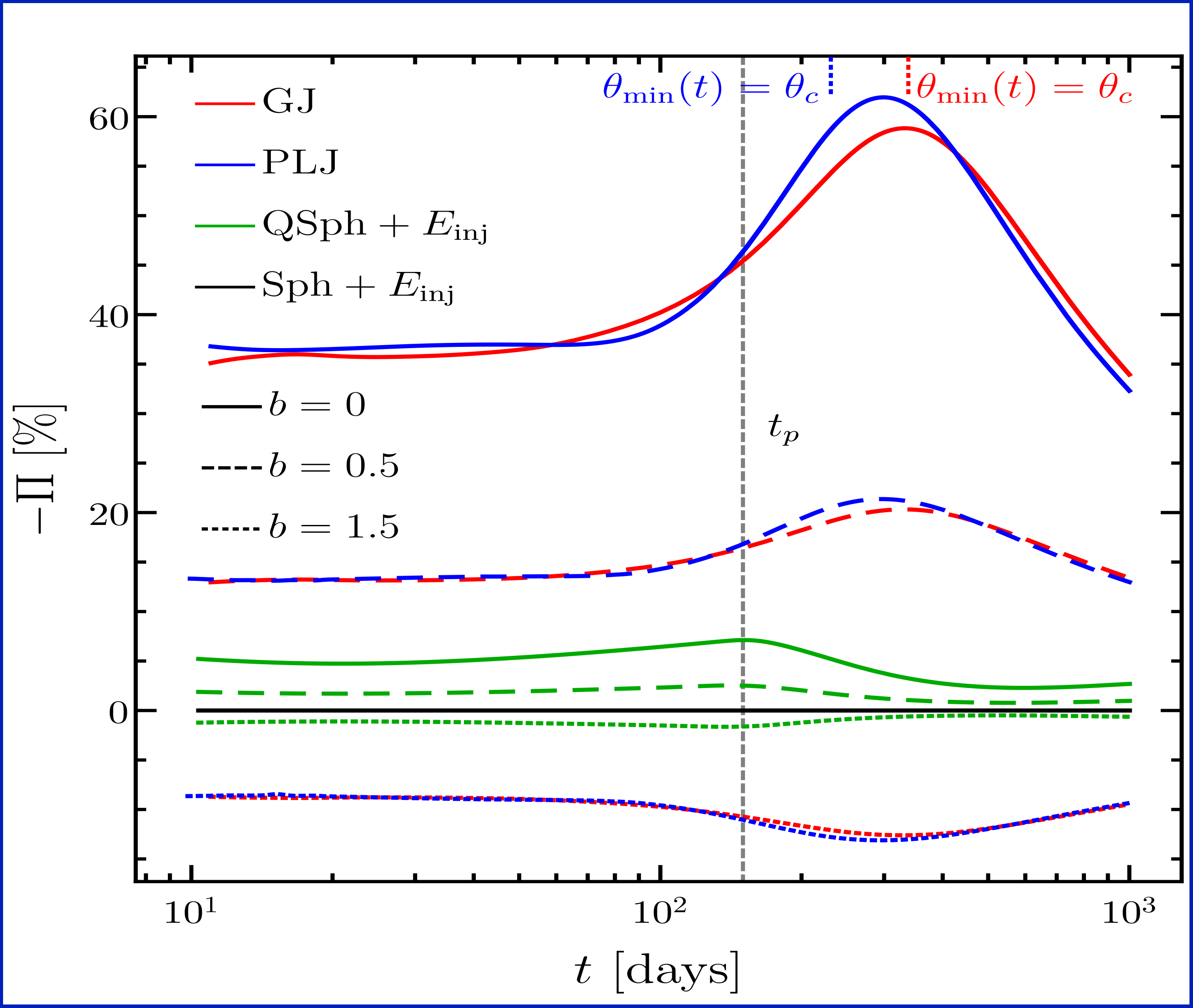
Afterglow Images: GJ, PLJ





Linear Polarization

- Assuming a shock-produce B-field with $b \equiv 2\langle B_{\parallel}^2 \rangle / \langle B_{\perp}^2 \rangle$



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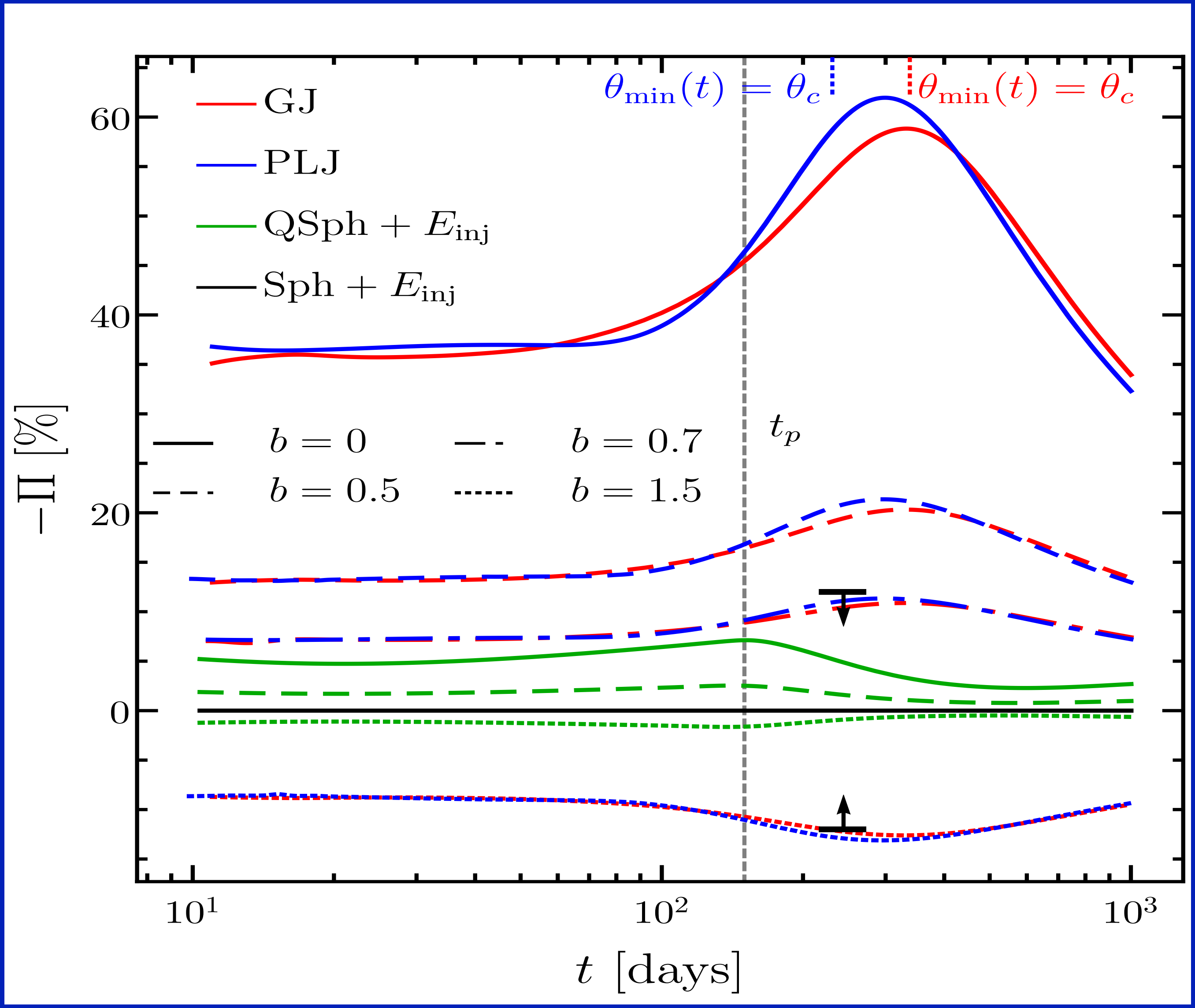


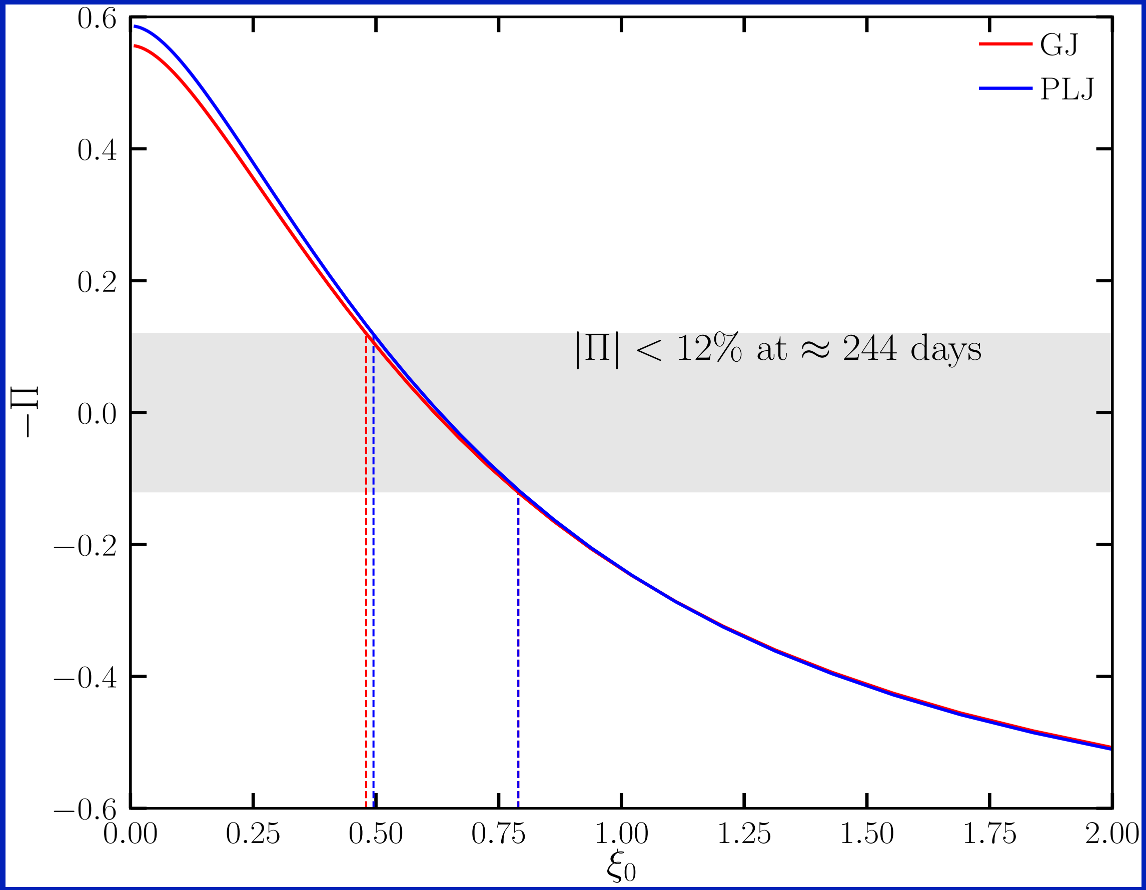
$0.7 \lesssim b \lesssim 1.5$ for jet models

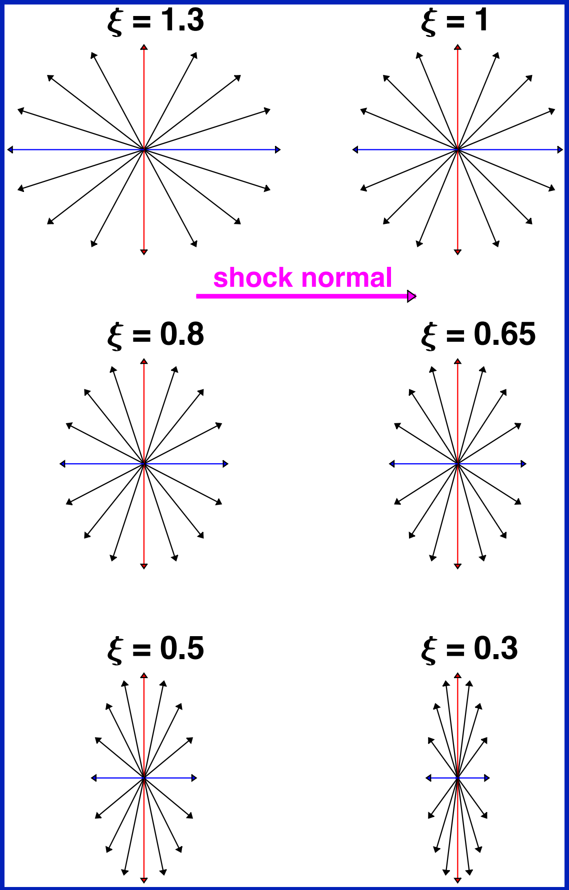
New: upper limit on linear pol. @ 2.8 GHz (Corsi + 2018)

Linear Polarization (Gill & JG 2019)

More realistic assumptions ⇒ B-field in collisionless shocks:

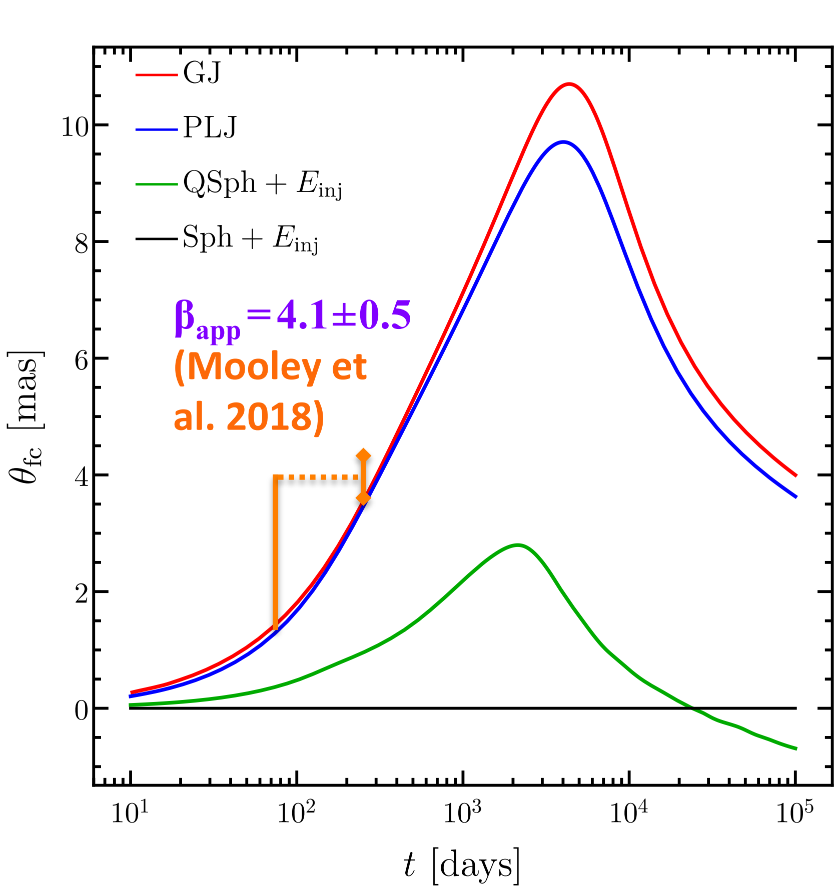
- 2D emitting shell → 3D emitting volume (local BM76 radial profile)
- B-field evolution by faster radial expansion: $L'_r / L'_{\theta,\varphi} \propto \chi^{(7-2k)/(8-2k)}$
  B-field isotropic in 3D with $B'_r \rightarrow \xi B'_r$ (Sari 1999); $\xi = \xi_0 \chi^{(7-2k)/(8-2k)}$

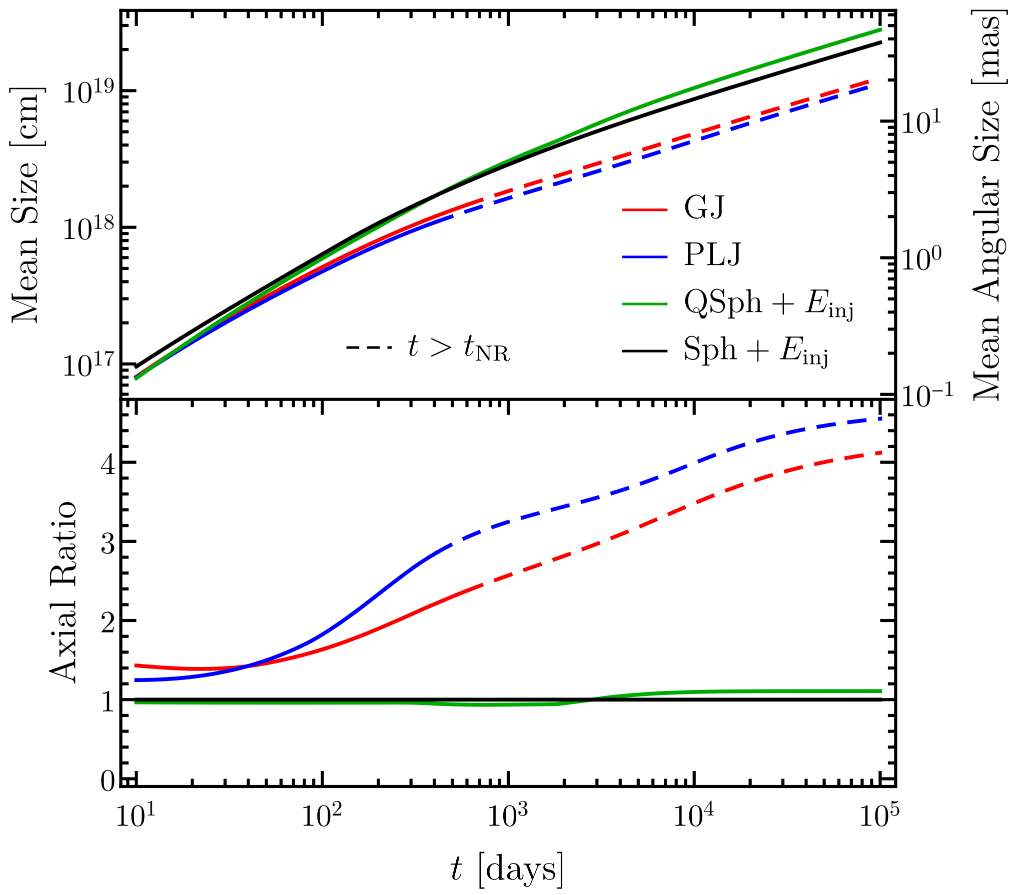




Afterglow Images: flux centroid, size, shape

- The flux centroid motion: a potentially powerful diagnostic
- It may be hard to tell apart models based on the image size alone, but a much higher axis-ratio is expected for jet models





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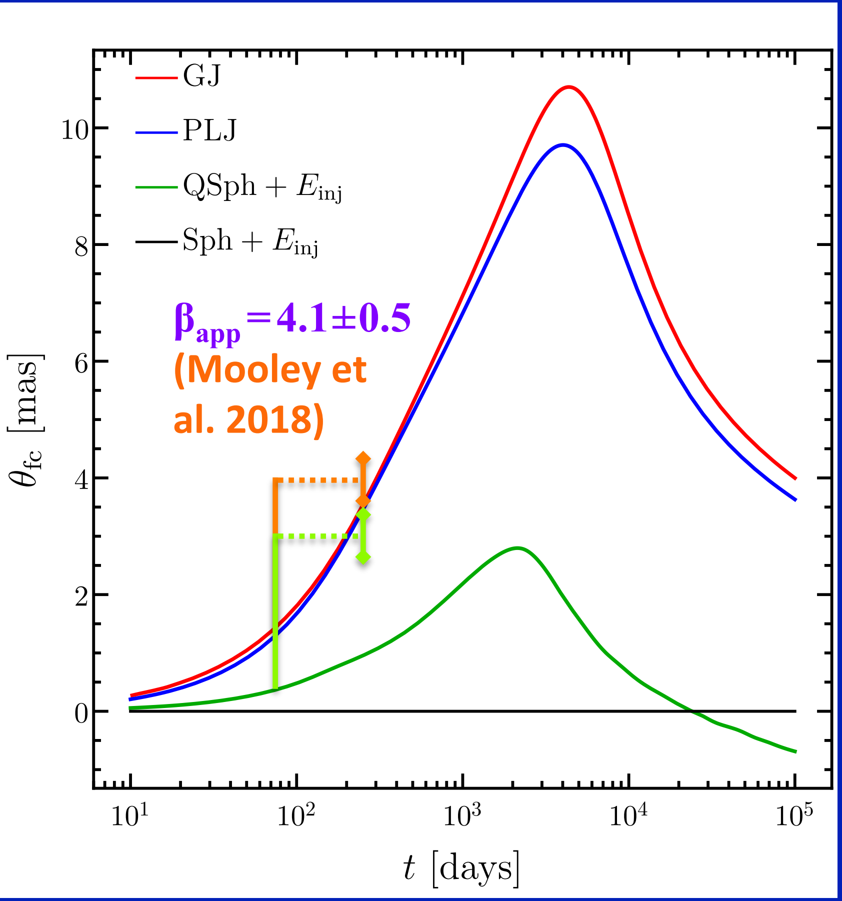

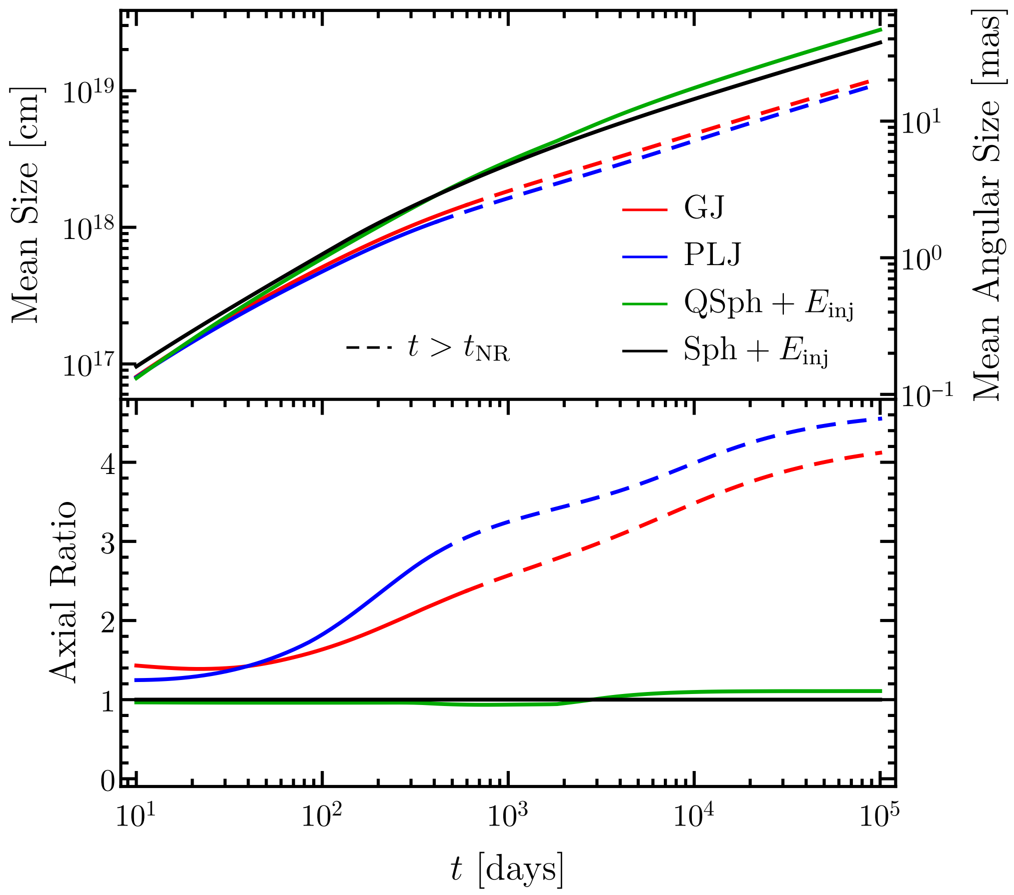

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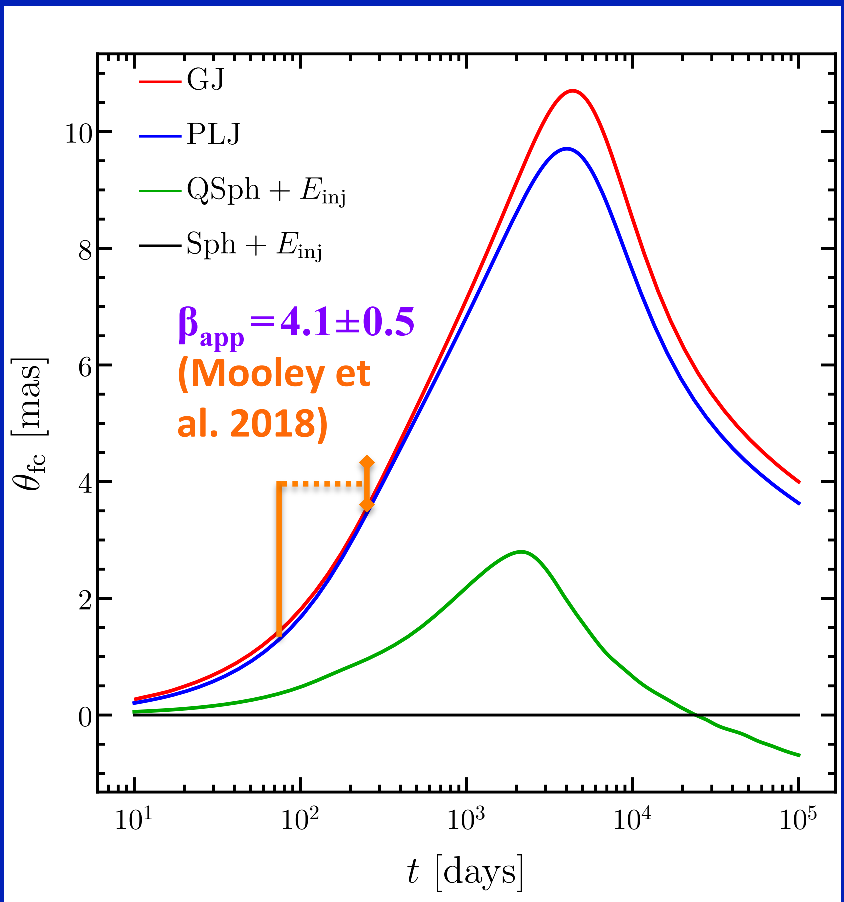
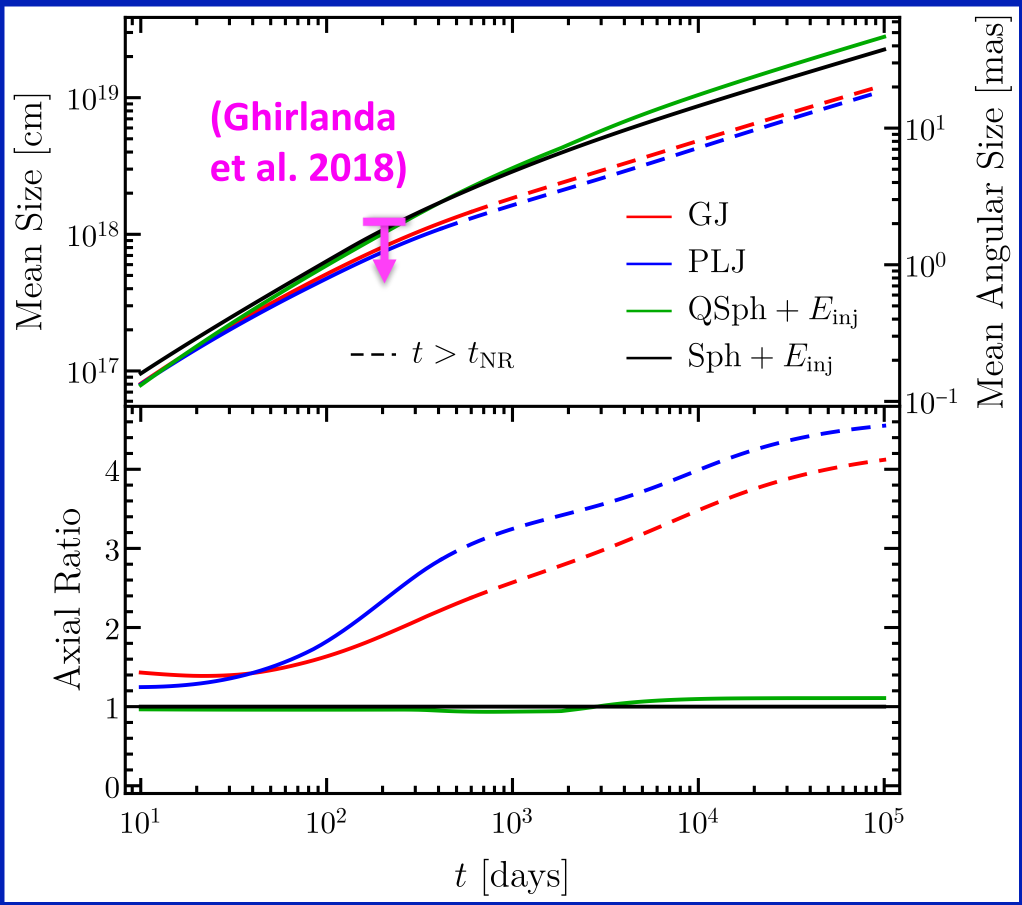
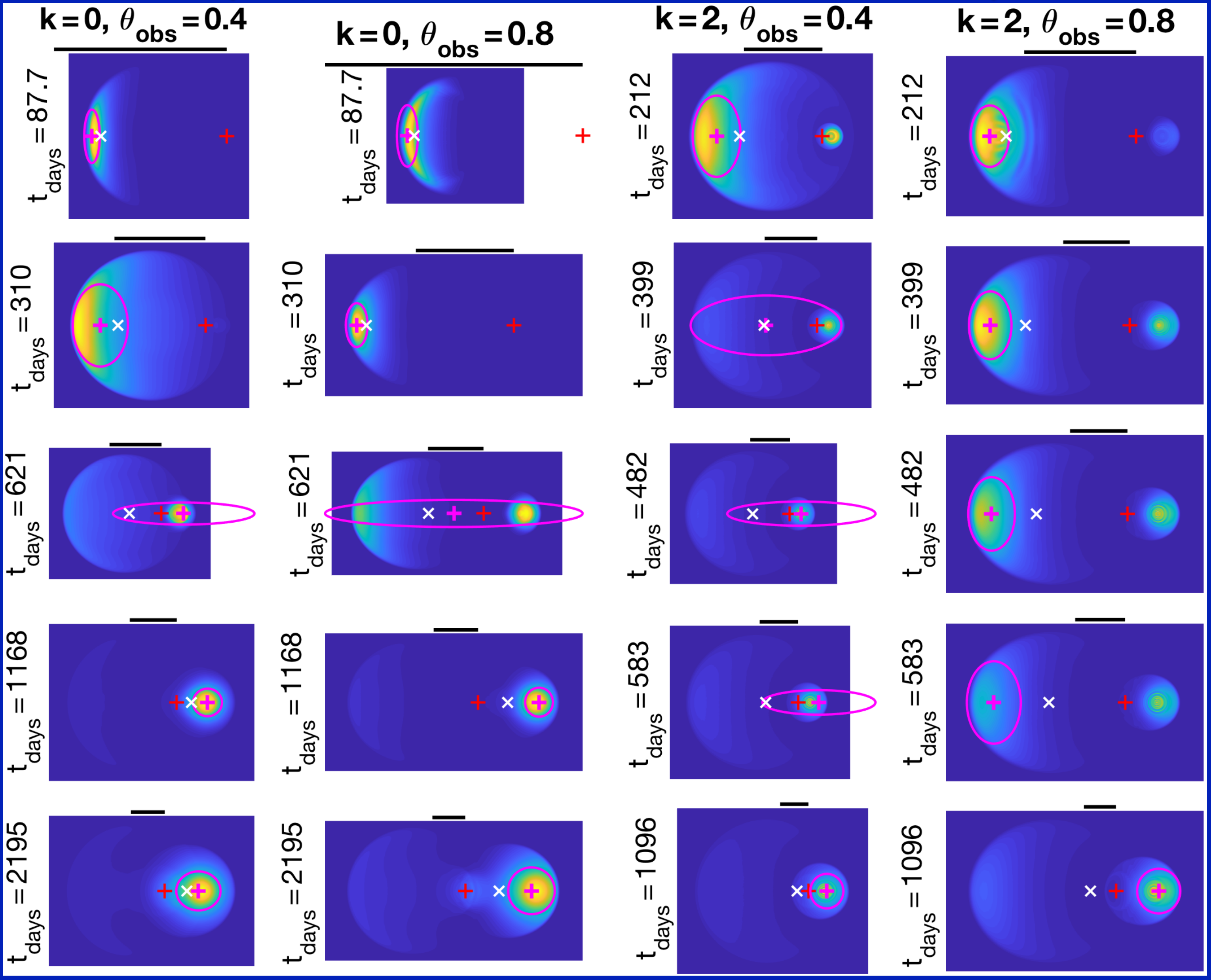




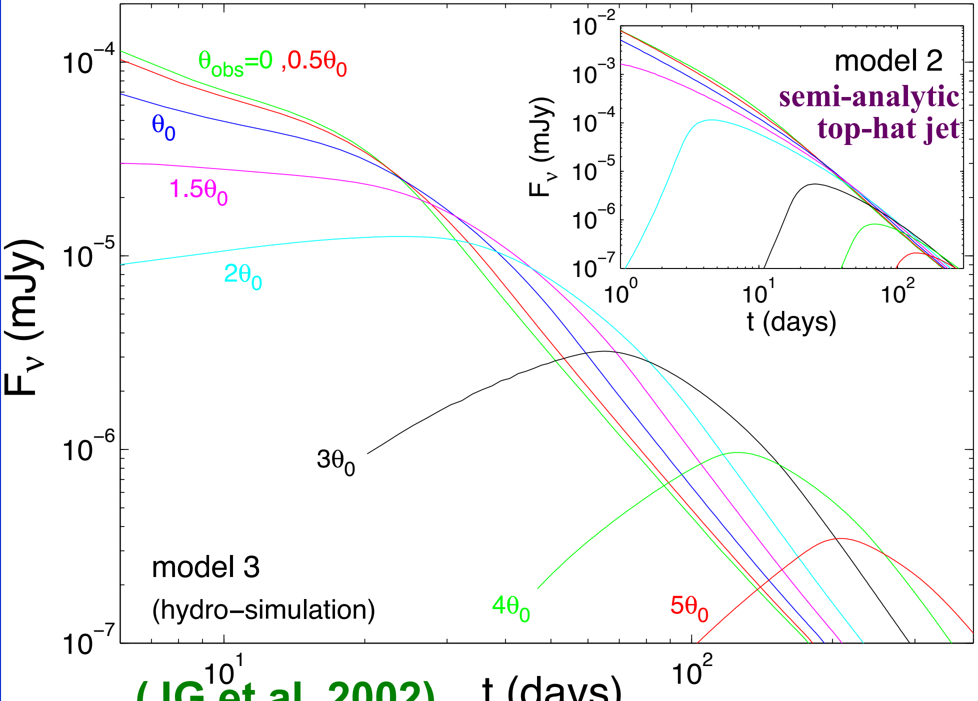
Afterglow Images: uniform jet simulations

(JG, De Colle & Remirez-Ruiz 2018)

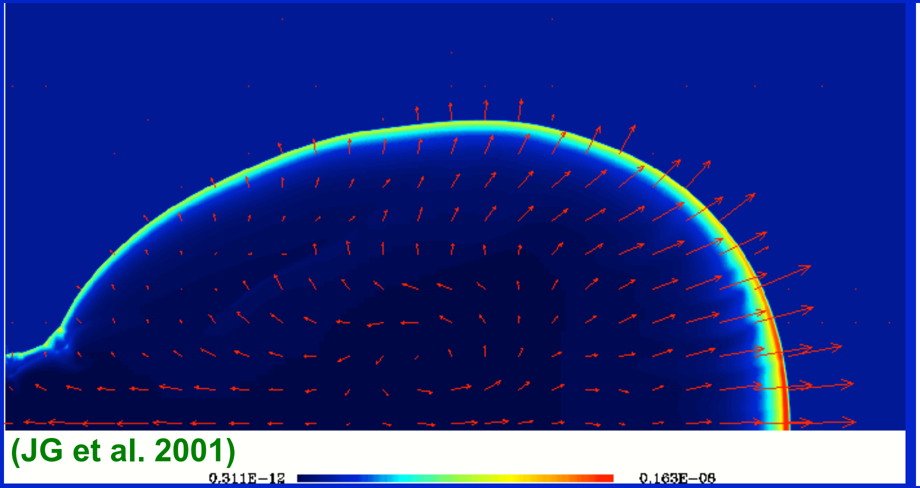


Off-Axis Afterglow Lightcurves: Top-Hat

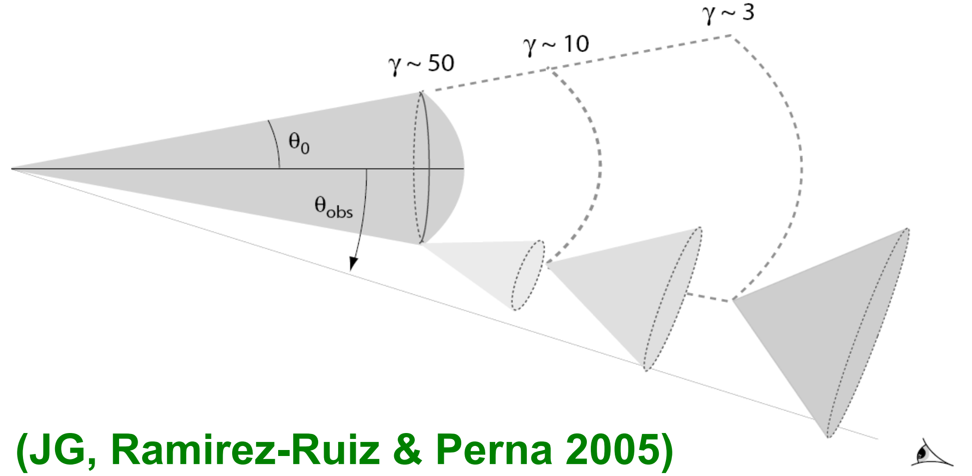
- The emission is initially strongly beamed away from our L.o.S
- F_ν rises as beaming cone widens
- When beaming cone reaches LoS F_ν peaks & approaches on-axis F_ν
- The rise is much more gradual for hydrodynamic simulations due to slower matter at the jet's sides with non-radial velocities



(JG et al. 2002)



(JG et al. 2001)



(JG, Ramirez-Ruiz & Perna 2005)

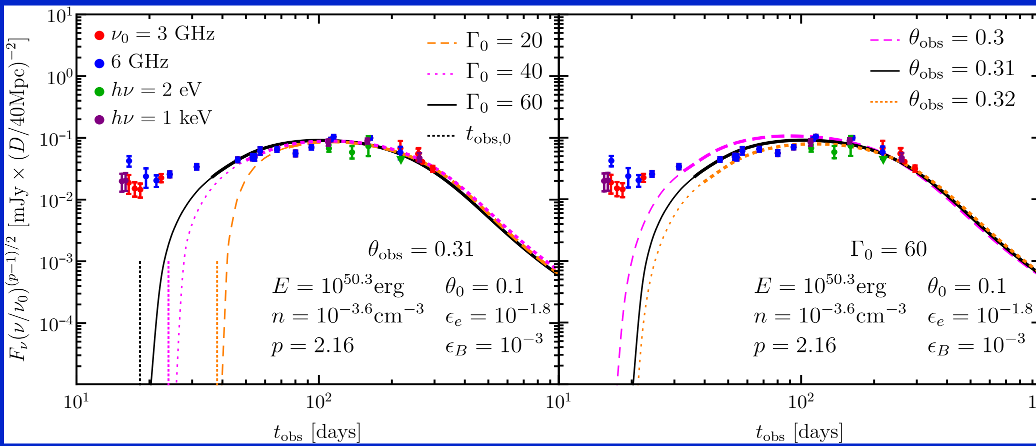
GW170817 afterglow: Top-Hat Jet? (Gill & JG 19)

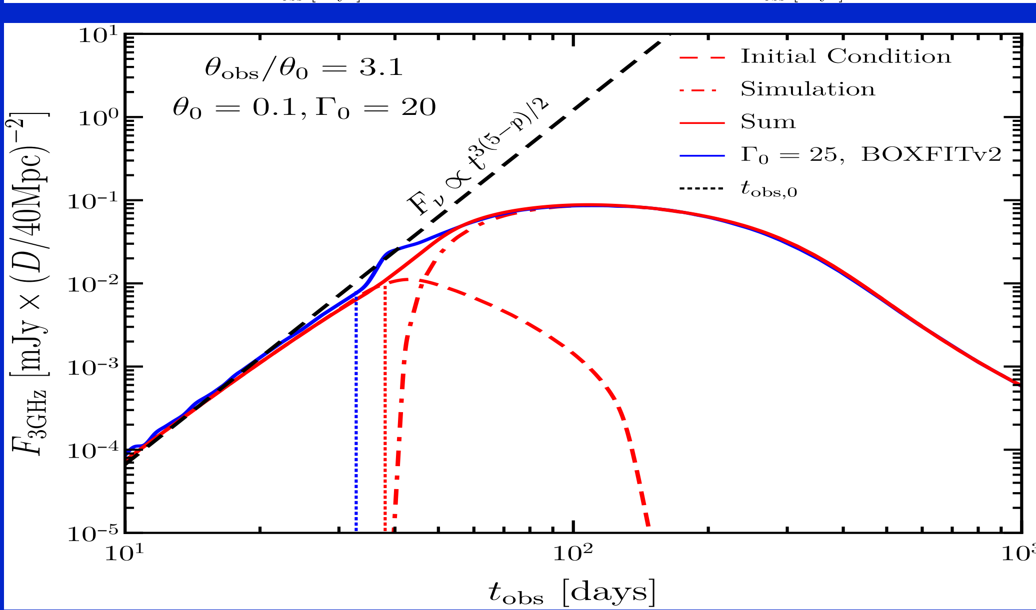
- In core-dominated jets angular structure hardly affects $t \gtrsim t_{peak}$ emission

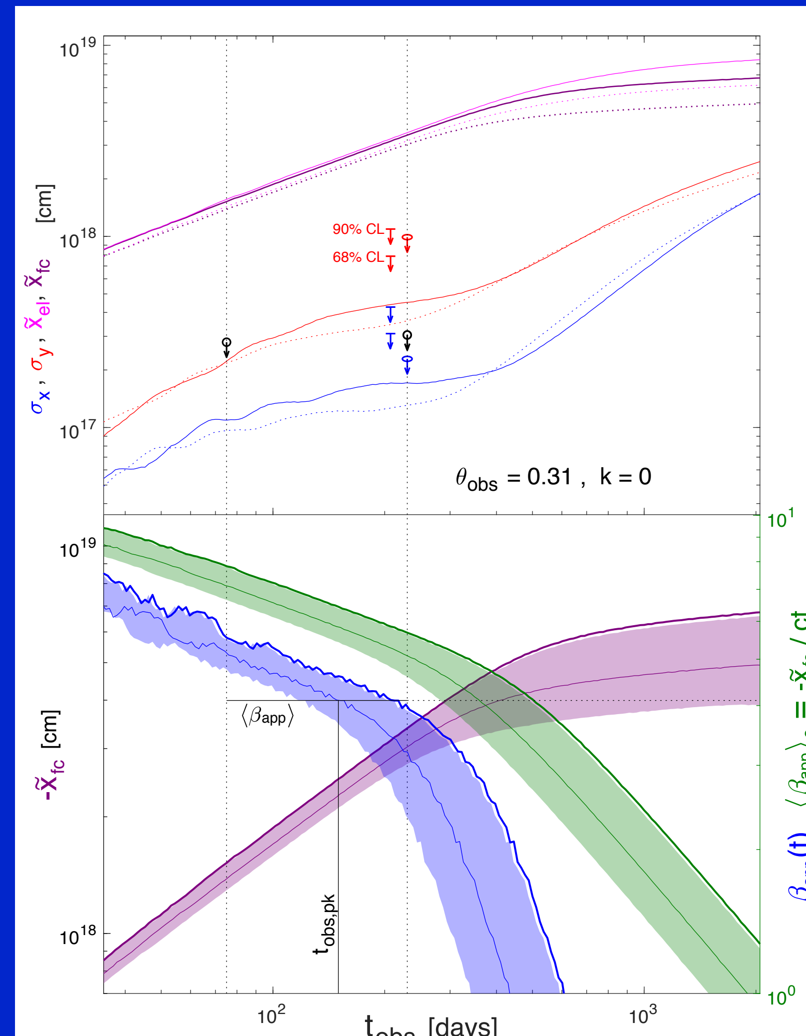
$$t_\theta = \frac{R_\gamma}{c}[1 - \cos(\Delta\theta)] \approx \frac{R_\gamma}{2c}\Delta\theta^2$$

$$t_{obs,0} = \frac{R_0}{c}[1 - \cos(\Delta\theta)] \approx \frac{R_0}{2c}\Delta\theta^2$$

Image Size & Flux Centroid

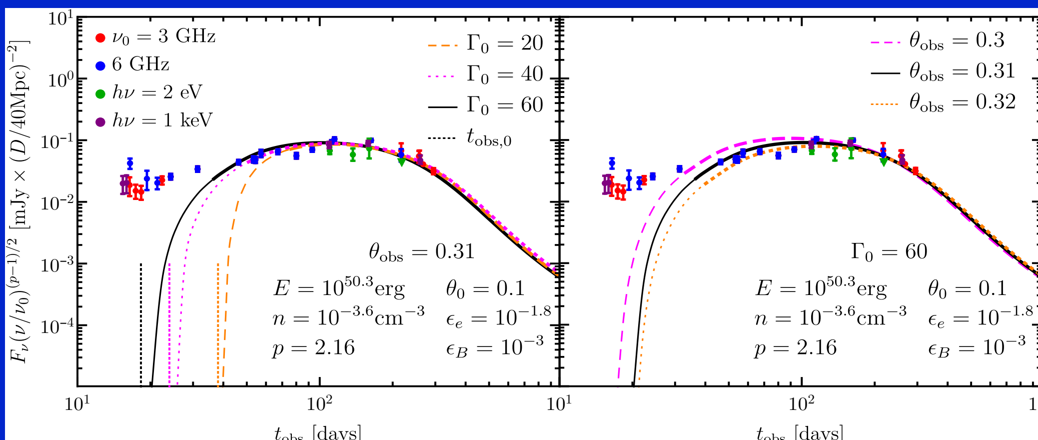


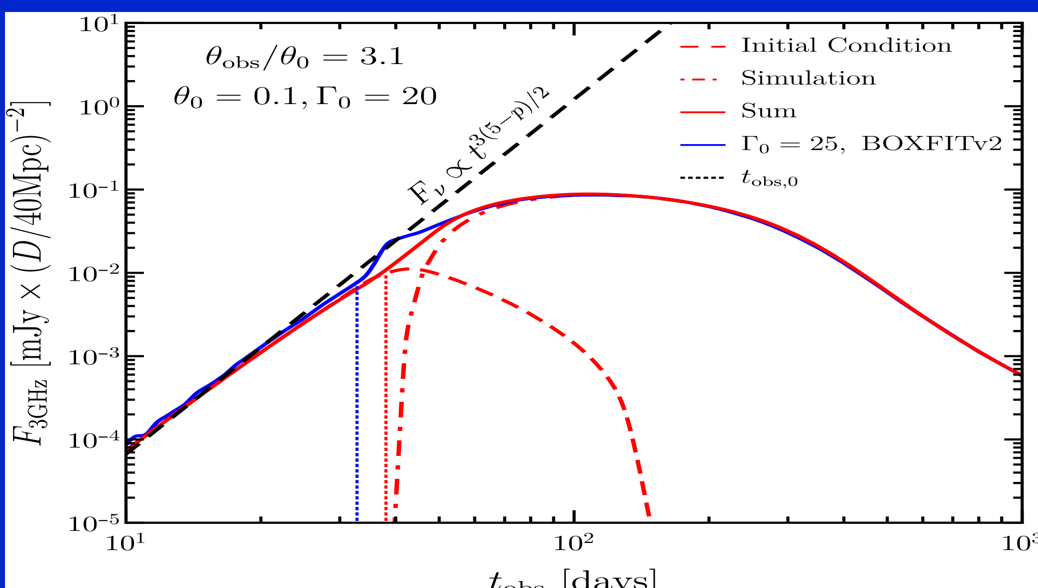


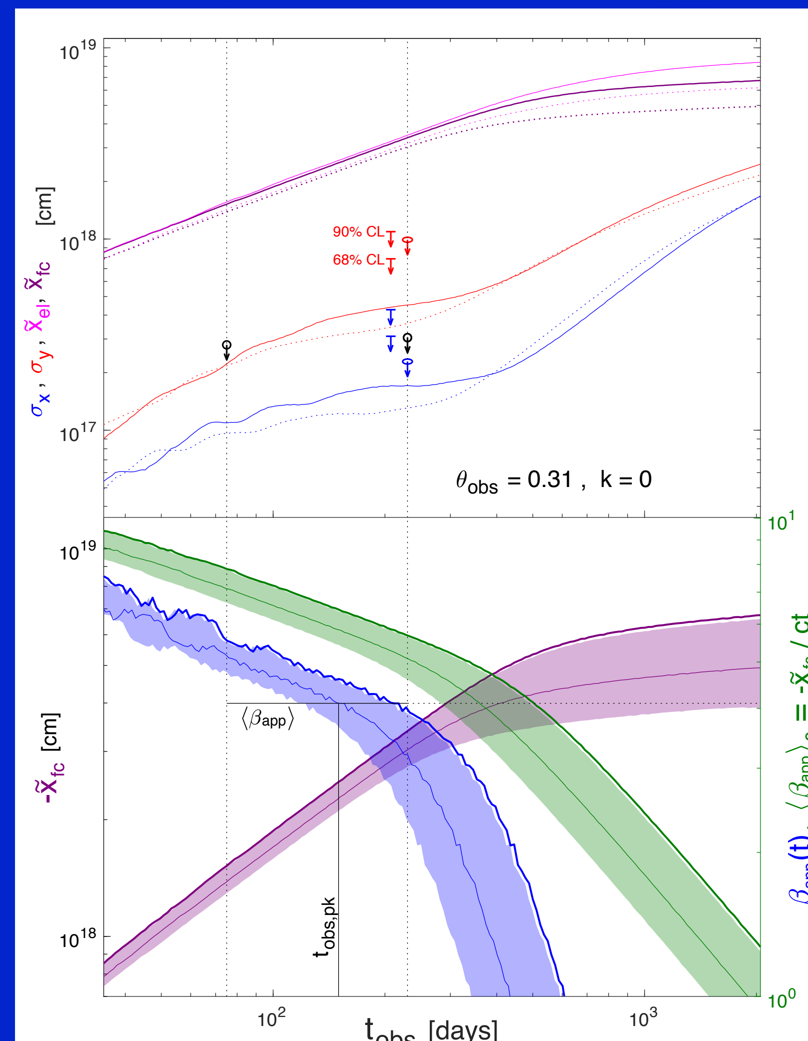


GW170817 afterglow: Top-Hat Jet? (Gill & JG 19)

- In core-dominated jets angular structure hardly affects $t \gtrsim t_{peak}$ emission
- Both initial angular structure & dynamics affect the early afterglow rise
- The initial angular structure strongly affects the prompt GRB at $\theta_{obs} > \theta_c$







Conclusions on GW170817/GRB170817A:

- It is a unique event with a wide range of implications
- Merger Remnant: BH or HMNS → BH ⇒ $M_{max} \lesssim 2.17 M_\odot$
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